

### The Galactic Center Environment and Dark Matter Annihilation Tim Linden

CCAPP Postdoctoral Fellow Center for Cosmology and Astro-Particle Physics The Ohio State University



The Ohio State University

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### The Central Molecular Zone

- 400 pc x 80 pc
- 10<sup>7</sup> M<sub>o</sub> of gas in Molecular Clouds
- Conditions similar to nearby starburst galaxies





- Molecular Gas clouds in the Central Molecular Zone are hot (~50-100K)
- Indicative of heating by a significant cosmic-ray population confined in the central molecular zone. (Yusef-Zadeh et al. 2013)

### What Generates these Cosmic-Rays?

The Galactic center region is known to contain nearly every known cosmic-ray acceleration mechanism.

- 1.) Supernovae
- 2.) Pulsars
- 3.) Sgr A\*
- 4.) Reacceleration
- 5.) Dark Matter Annihilation?





#### The Galactic Center Supernovae

Multiwavelength observations indicate that the Galactic Center is a dense star-forming environment.

3-20% of the total Galactic Star Formation Rate is contained within the Central Molecular Zone.

2-4% - ISOGAL Survey Immer et al. (2012)
2.5-5% - Young Stellar Objects Yusef-Zadeh et al. (2009)
5-10% - Infrared Flux Longmore et al. (2013)
10-20% - Wolf-Rayet Stars Rosslowe & Crowther (2014)
2% - Far-IR Flux Thompson et al. (2007)
2.5-6% - SN1a Schanne et al. (2007)



Arches Cluster Ə<sub>GC</sub>=0.25°, Age~2 Myr

### **Galactic Center Pulsars**

Chandra Observes > 9000 point sources from the inner 1° x 0.5°

The Galactic Center is expected to host a significant population of both young pulsars (due to its high SFR), and millisecond pulsars (in part from the disruption of Globular Clusters).

Over the lifetime of a young (recycled) pulsar, ~10<sup>50</sup> erg of energy our released, primarily in the form of relativistic  $e^+e^-$  pairs.

### The Sgr A\* Source

#### see slides by: R. Chaves, R. Parsons



HESS has detected diffuse gamma-ray emission at energies ~100 TeV.

This is not observed in even the youngest supernova remnants.

The emission profile is indicative of diffusion from the central BH.



### Dark Matter Annihilation?

WIMPs are currently among the most well-motivated dark matter models.

WIMP annihilation naturally produces a significant cosmic-ray (and gamma-ray) flux.





Dark Matter structure simulations uniformly predict that the GC is the brightest source of WIMP annihilations.

Standard scenarios predict the flux from the GC exceeds dSphs by a factor of ~100 – 1000.

### Reacceleration

More than 80 filamentary structures identified in the central 2° x 1°.

The filaments are observed as highly polarized, hard-spectrum synchrotron sources – indicative of strongly ordered magnetic fields and hard injected electron spectra.

The best astrophysical explanation involves significant re-acceleration via magnetic reconnection (Lesch & Riech 1992, Lieb et al. (2004).





Yusef-Zadeh et al. (2004)

#### Non-Thermal Emission (Observables) Fermi Bubbles GeV Excess



#### WMAP/PLANCK Haze

#### Integral 511 keV Excess

#### Non-Thermal Emission (Observables)

Fermi Bubbles

GeV Excess



#### The photon excesses extend very far from the central molecular region!

#### This:

(a) Indicates the relative power of Galactic center accelerators, compared to the Galactic plane.
(b) Provides a large field of view for studies of GC emission.
(c) Implies that propagation is important!



#### Start with a source of relativistic cosmic-rays

### Start with a source of relativistic cosmic-rays cosmic rays propagate $\frac{\partial \psi}{\partial t} = q(\vec{r}, p) + \vec{\nabla} \cdot (D_{xx} \vec{\nabla} \psi - \vec{V} \psi) + \frac{\partial}{\partial p} p^2 D_{pp} \frac{\partial}{\partial p} \frac{1}{p^2} \psi - \frac{\partial}{\partial p} \left[ \dot{p} \psi - \frac{p}{3} (\vec{\nabla} \cdot \vec{V}) \psi \right]$ **Solved Numerically:** e.g. Galprop see slides by T. Porter

#### **Start with a source of relativistic cosmic-rays**



#### cosmic rays propagate



Solved Numerically: e.g. Galprop

see slides by T. Porter





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e.g. Galprop

see slides by T. Porter









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Approximately half of this emission is produced along the line of sight towards the GC, and thus we approximate the total gamma-ray luminosity of the central one degree to be  $5 \times 10^{36} \text{ erg s}^{-1}$ 

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A Supernovae produces ~10<sup>51</sup> erg of energy.

~10% to CR protons.

Assuming 1 Galactic center SN every 250 years (10% the Galactic Rate), this provides an energy flux of  $1.3 \times 10^{40}$  erg s<sup>-1</sup>.

If these cosmic-rays are trapped for 10 kyr in a 100 pc box ( $D_0 = 5 \times 10^{28} \text{ cm}^2 \text{ s-1}$ ), filled with Hydrogen gas at density 100 cm<sup>-2</sup>, this will produce a total gamma-ray emission:



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Sgr A\* is expected to produce a tidal disruption event every ~10<sup>5</sup> yr, producing a timeaveraged energy output of 2 x  $10^{39}$  erg s<sup>-1</sup>.

If these CRs are primarily leptonic, and the electrons remain trapped in a region with a 40 eV cm<sup>-3</sup> ISRF and a 200  $\mu$ G magnetic field the gamma-ray flux from inverse Compton scattering is:



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MSPs observed in the galactic field are fit by a population with a mean gamma-ray flux of  $3 \times 10^{34}$  erg s<sup>-1</sup>. (Hooper & Mohlabeng 2015)

Given the population of 129 MSPs among 124 globular clusters (with a total stellar mass ~5 x 10<sup>7</sup> M<sub>o</sub>). For the 1 x 10<sup>9</sup> M<sub>o</sub> of stars formed in the inner degree of the Milky Way, we get:



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#### **Dark Matter**

For a 35 GeV dark matter particle annihilating at the thermal cross-section to bb, and a slightly adiabatically contracted r<sup>-1.35</sup> density profile.

The dark matter annihilation rate is 8.6 x 10<sup>38</sup> ann s<sup>-1</sup>, which produces a gamma-ray flux of:



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# Conclusion: Every Model is Correct

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# Is this (theorist) heaven? or is this hell?



### **Modeling the Galactic Center**





#### Two Analyses of the Gamma-Ray Excess



#### **INNER GALAXY**

- Mask galactic plane (e.g. |b| > 1°), and consider 40° x 40° box
- Bright point sources masked at 2°
- Use likelihood analysis, allowing the diffuse templates to float in each energy bin
- Background systematics controlled

#### **GALACTIC CENTER**

- Box around the GC (10° x 10°)
- Include and model all point sources
- Use likelihood analysis to calculate the spectrum and intensity of each source
- Bright Signal

### The Galactic Center Excess

Utilizing different models for removing astrophysical and point source foregrounds. Multiple studies have consistently observed a gamma-ray excess.

> Goodenough & Hooper (2009, 0910.2998) Hooper & Goodenough (2010, 1010.2752) Hooper & Linden (2011, 1110.0006) Abazajian & Kaplinghat (2012, 1207.6047) Gordon & Macias (2013, 1306.5725) Gordon & Macias (2013, 1312.6671) Abazajian et al. (2014, 1402.4090) Daylan et al. (2014, 1402.6703) Calore et al. (2014, 1409.0042) Abazajian et al. (2014, 1410.6168) Bartels et al. (2015, 1506.05104) Lee et al. (2015, 1506.05124) Gaggero et al. (2015, 1507.06129) Carlson et al. (2015, 1510.04698) The Fermi-LAT Collaboration (2015, 1511.02938) Yang & Aharonian (2016, 1602.06764) Carlson et al. (2016, 1603.06584) Linden et al. (2016, 1604.01026) Horiuichi et al. (2016, 1604.01402)



#### Daylan <u>et al. (2014)</u>

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### **Observational Results**



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The excess has an unusual spectrum - highly peaked at an energy of ~2 GeV.

This spectrum is significantly harder than expected from astrophysical diffuse emission.

### **Observational Results**





The GeV excess spherically symmetric, and is statistically significant from  $0.1^{\circ} - 10^{\circ}$  from the Galactic Center.

#### **Observational Results**

<u>These are the three resilient features of the GeV Excess</u>:
1.) Hard Gamma-Ray Spectrum peaking at ~2 GeV
2.) Spherically Symmetric Emission Morphology
3.) Extension to >10° from the GC.

#### How could we model this with:

- 1.) Dark Matter annihilation
- 2.) Millisecond Pulsars
- 3.) Leptonic Outbursts from Sgr A\*
- 4.) Changes in Diffuse Emission Modeling

#### Dark Matter Model Fitting?









### Particle Physics Models Exist...

Chan (1607.02246) Jia (1607.00737) Barrau et al. (1606.08031) Huang et al. (1605.09018) Cui et al. (1605.08138) Krauss et al. (1605.05327) Kumar et al. (1605.00611) Biswas et al. (1604.06566) Sage et al. (1604.04589) Choquette et al. (1604.01039) Cuoco et al. (1603.08228) Chao et al. (1602.05192) Horiuchi et al. (1602.04788) Hektor et al. (1602.00004) Freytsis et al. (1601.07556) Kim et al. (1601.05089) Huang et al. (1512.08992) Kulkami et al. (1512.06836) Tang et al. (1512.02899) Cox et al. (1512.00471) Cai et al. (1511.09247) Agrawal et al. (1511.06293) Duerr et al. (1510.07562) Drozd et al. (1510.07053) Arcadi et al. (1510.02297) Williams (1510.00714) Cai & Spray (1509.08481) Freese et al. (1509.05076) Bhattacharya et al. (1509.03665) Algeri et al. (1509.01010) Fox & Tucker-Smith (1509.00499) Dutta et al. (1509.05989) Liu et al. (1508.05716) Berlin et al. (1508.05390) Fan et al. (1507.06993) Hektor et al. (1507.05096) Achterbeg et al. (1507.04644) Biswas et al. (1507.04543)

Butter et al. (1507.02288) Mondal et al. (1507.01793) Cao et al. (1506.06471) Banik et al. (1506.05665) lpek (1505.07826) Buchmueller et al. (1505.07826) Balazs et al. (1505.06758) Medina (1505.05565) Kim et al. (1505.04620) Ko et al. (1504.06944) Ko & Tang (1504.03908) Ghorbani & Ghorbani (1504.03610) Fortes et al. (1503.08220) Cline et al. (1503.08213) Rajaraman et al. (1503.05919) Bi et al. (1503.03749) Kopp et al. (1503.02669) Elor et al. (1503.01773) Gherghetta et al. (1502.07173) Berlin et al. (1502.06000) Achterberg et al. (1502.05703) Modak et al. (1502.05682) Guo et al. (1502.00508) Chen & Nomura (1501.07413) Kozaczuk & Martin (1501.07275) Berlin et al. (1501.03496) Kaplinghat et al. (1501.03507) Alves et al. (1501.03490) Biswas et al. (1501.02666) Biswas et al. (1501.02666) Ghorbani & Ghorbani (1501.00206) Cerdeno et al. (1501.01296) Liu et al. (1412.1485) Hooper (1411.4079) Arcadi et al. (1411.2985) Cheung et al. (1411.2619) Agrawal et al. (1411.2592) Kile et al. (1411.1407)

Buckley et al. (1410.6497) Heikinheimo & Spethmann (1410.4842) Freytsis et al. (1410.3818) Yu et al. (1410.3347) Cao et al. (1410.3239) Guo et al. (1409.7864) Yu (1409.3227) Cahill-Rowley et al. (1409.1573) Banik & Majumdar (1408.5795) Bell et al. (1408.5142) Ghorbani (1408.4929) Okada & Seto (1408.2583) Frank & Mondal (1408.2223) Baek et al. (1407.6588) Tang (1407.5492) Balazs & Li (1407.0174) Huang et al. (1407.0038) McDermott (1406.6408) Cheung et al. (1406.6372) Arina et al. (1406.5542) Chang & Ng (1406.4601) Wang & Han (1406.3598) Cline et al. (1405.7691) Berlin et al. (1405.5204) Mondal & Basak (1405.4877) Martin et al. (1405.0272) Ghosh et al. (1405.0206) Abdullah et al. (1404.5503) Park & Tang (1404.5257) Cerdeno et al. (1404.2572) Izaguirre et al. (1404.2018) Agrawal et al. (1404.1373) Berlin et al. (1404.0022) Alves et al. (1403.5027) Finkbeiner & Weiner (1402.6671) Boehm et al. (1401.6458) Kopp et al. (1401.6457) Modak et al. (1312.7488)

Alves et al. (1312.5281) Fortes et al. (1312.2837) Banik et al. (1311.0126) Arhrib et al. (1310.0358) Kelso et al. (1308.6630) Kozaczuk et al. (1308.5705) Kumar (1308.4513) Demir et al. (1308.1203) Buckley et al. (1307.3561) Cline et al. (1306.4710) Cannoni et al. (1205.1709) An et al. (1110.1366) Buckley et al. (1106.3583) Boucenna et al. (1106.3368) Ellis et al. (1106.0768) Cheung et al. (1104.5329) Marshall et al. (1102.0492) Abada et al. (1101.0365) Tytgat (1012.0576) Logan (1010.4214) Barger et al. (1008.1796) Raklev et al. (0911.1986)

## **Comparison to Dwarf Constraints**

#### Ackermann et al. (2015)



However, uncertainties in the dark matter density profile can easily resolve this tension.

credit: Kev Abazajian (2015)

Constraints from dSphs are statistically in 1-2 $\sigma$  tension with the GC excess.



### Millisecond Pulsar Fits

 The peak of the MSP energy spectrum matches the peak of the GeV excess





 MSPs are thought to be overabundant in dense star-forming regions like the Galactic Center

see slides by: Christoph Weniger and Emma Storm
# **Millisecond Pulsar Fits**

Bartels et al. (2015)



Lee et al. (2015)

 Recent analyses of hot-spots and cold spots in the GC region find evidence for the presence of a population of subthreshold point sources.

see slides by Christoph Weniger

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# Millisecond Pulsar Fits



 However, these residuals are found once an extremely smooth diffuse emission model is subtracted - it remains to be seen whether the residuals are resilient to diffuse model changes.

see slides by Christoph Weniger

Ajello et al. (2015)

## **Millisecond Pulsar Fits**

- Millisecond pulsars in the Galactic center must be categorically dimmer and more numerous than those in the Galactic plane or in globular clusters
  - Using a luminosity function similar to that of the plane produces too many detectable MSPs (Hooper et al. 2013)
  - A comparison to the LMXB population in Globular Clusters indicates that MSPs can only account for 1-5% of the excess (Hooper et al. 2015)
  - MSPs from disrupted globular clusters could potentially be more numerous, but would need  $10^9 M_o$  of disrupted globular cluster material to generate the excess. (Hooper & TL 2016)

#### Calore et al. (2015) Fortunately the Pulsar Hypothesis is Testable

- Radio Observations with GBT targeted at gamma-ray hotspots would be expected to find ~5-10 MSPs with a 200 hr commitment.
- Fortunately, SKA observations are likely to conclusively find MSPs in the GC, or rule out this scenario.





### Leptonic Outbursts

The Galactic center is unlikely to be in steady state (e.g. Fermi bubbles).

An outburst of leptonic origin can produce the gamma-ray excess, but only if the injected electron spectrum is extremely hard (compared to observed blazar spectra).



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Petrovic et al. (2014, 1405.7928) Cholis et al. (2015, 1506.05119)

# An Excess Compared to What?

Cosmic-Ray Propagation Codes (e.g. Galprop), generally utilize a cosmicray injection rate at the Galactic center that is identically 0.

These models were not produced to study the very center of the Galaxy!





Results from these cosmic-ray propagation codes are used in many analyses of the Galactic center region.

> Carlson et al. (2016a, 2016b) 1510.04698 1603.06584

# **The Solution**

**Solution:** Add a new cosmic-ray injection morphology tracing the molecular gas density.

**Observationally Resilient:** Several tracers of molecular gas are sensitive to the galactic center region.

**Theoretically Motivated:** Molecular Gas is the seed of star formation, the Schmidt Law gives

 $\Sigma_{\rm SFR} \propto \Sigma_{\rm Gas}^{1.4\pm.15}$ 

Specifically we inject a fraction of cosmic-rays (0 <  $f_{H2}$  < 1) following:

$$\mathbf{Q}_{\mathrm{CR}}(\vec{r}) \propto \begin{cases} 0 & \rho_{\mathrm{H2}} \leq \rho_s \\ \rho_{\mathrm{H2}}^{n_s} & \rho_{\mathrm{H2}} > \rho_s \end{cases}$$

1510.04698

# The Solution



Two features leap out immediately:

1.) Spiral Arms

2.) A bright bar in the Galactic Center

# The Solution



Adds a new, and significant, cosmic-ray injection component, in particular near the Galactic Center.

The cosmic-ray injection rate now matches observational constraints.

# Simulations!

Add the new cosmic-ray injection models into Galprop to produce a new steady-state cosmic-ray distribution.

Parameter	Units	Canonical	Mod A	Description
$D_0$	$\rm cm^2~s^{-1}$	$7.2\times10^{28}$	$5.0\times10^{28}$	Diffusion constant at $\mathcal{R} = 4$ GV
δ	-	0.33	0.33	Index of diffusion constant energy dependence
$z_{\rm halo}$	kpc	3	4	Half-height of diffusion halo
$R_{halo}$	kpc	20	20	Radius diffusion halo
$v_a$	$\rm km \ s^{-1}$	35	32.7	Alfvén velocity
dv/dz	$\rm km~s^{-1}~kpc^{-1}$	0	50	Vertical convection gradient
$\alpha_{\rm p}$	-	1.88 (2.39)	1.88 (2.47)	p injection index below (above) $\mathcal{R} = 11.5 \text{ GV}$
$\alpha_{e}$	-	1.6(2.42)	1.6(2.43)	$e^-$ injection index below (above) $\mathcal{R} = 2$ GV
Source	-	SNR	SNR	Distribution of $(1 - f_{H2})$ primary sources <sup>*</sup>
$f_{\rm H2}$	-	.20	N/A	Fraction of sources in star formation model*
$n_s$	-	1.5	N/A	Schmidt Index <sup>*</sup>
$\rho_c$	$\rm cm^{-3}$	0.1	N/A	Critical H <sub>2</sub> density for star formation <sup>*</sup>
$B_0$	μG	7.2	9.0	Local $(r = R_{\odot})$ magnetic field strength
$r_B, z_B$	kpc	5, 1	5, 2	Scaling radius and height for magnetic field
ISRF	_	(1.0, .86, .86)	(1.0, .86, .86)	Relative CMB, Optical, FIR density
dx, dy	kpc	0.5, 0.5	1 (2D)	x, y (3D) or radial (2D) cosmic-ray grid spacing
dz	kpc	0.125	.1	z-axis cosmic-ray grid spacing



# A Better fit to the Gamma-Ray Sky

1.) Adding a cosmic-ray injection component tracing  $f_{H2}$  improves the full-sky fit to the gamma-ray data.

2.) The best fit value over the full sky is  $f_{H2} = 0.25$ 



3.) Technique will become more powerful with the introduction of 3D gas and dust maps in the near future.

# **Application to the Galactic Center**



Best Angular Resolution Cut 10° x 10° ROI



# **Effect on the GC Excess**



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Increasing the value of  $f_{H2}$  decreases the intensity of the gamma-ray excess.

However, the best global fit is  $f_{H2} = 0.1$ , with a GC excess intensity that decreases by only ~30%.



The morphology of the excess is also degenerate with  $f_{\rm H2}$ .

As  $f_{H2}$  is increased, the best-fit morphology becomes stretched perpendicular to the galactic plane.

However, marginalized over all values of  $f_{H2}$ , the standard NFW template is still consistent with the data.

# Analysis in the Galactic Center



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In this smaller region, the excess remains resilient to changes in diffuse emission modeling.

### **The Galactic Center Deficit?**





### **Advection and Convection in the Galactic Center**

Crocker et al. (2011) demonstrated that the break in the GC synchrotron spectrum is best fit in the regime with:

a.) Large Magnetic Fields b.) Large Convective Winds

Very different from typical Galprop diffusion scenario.



### The Low Energy Spectrum



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Applying strong convective winds to the diffuse emission model fixes the low-energy over subtraction.

The intensity of the excess near the spectral peak also increases, up to ~50% of its nominal value.

The model produces a significantly better fit to the gamma-ray sky dataset - and also coincides better with multi wavelength data.

### **A Similar Result with Different Techniques**





# Waxing Philosophical....



The lack of cosmic-ray injection in the GC should still be slightly disturbing. Especially when we try to answer the question: "excess compared to what?"

Our models indicate a degeneracy between cosmic-ray injection and the existence of a Galactic center excess template tracing an NFW profile. However, at present the best fit models still include a significant NFW component.

# Not a Conclusion Yet...

1.) The Galactic Center is a complex, but exciting environment. Several significant excesses are tied to the dynamics of the Galactic center environment.

2.) The GeV excess is a robust component of the Galactic center emission profile. At present, no models have successfully eliminated the excess.

3.) Improving our diffuse emission modeling is imperative to understanding the properties of the excess.

### **Stay Tuned!**

# Extra Slides

# **Effect on the Excess Spectrum**

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Changing the morphology of the excess has a significant effect on the spectrum of the gamma-ray excess.

The spectrum becomes extremely hard as f<sub>H2</sub> is increased, most likely indicating that the GCE template is picking up mismodeling of some residual.



# Dark Matter Annihilation?



Recently, observations by locco, Pato & Bertone (2015) have used stellar velocity measurements to directly measure the dark matter density in the Milky Way (to within 3 kpc of the GC).

Future measurements (employing Gaia data) will have the ability tosignificantly improve these measurements.locco, Pato & Bertone (2015)

### Galactic center excess is resilient....



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# The Sgr A\* Source

$\gamma(1 \text{ Lev} \geq E_{\gamma})[\times 10 \text{ ergs}]$	Mass [ $\times 10^{6} M_{\odot}$ ]		
	CS	CO	HCN
$1.54\pm0.06_{\rm stat}\pm0.23_{\rm sys}$	1.35	0.88	0.88
$1.21\pm0.05_{\rm stat}\pm0.18_{\rm sys}$	1.11	0.98	0.86
$2.05\pm0.08_{\rm stat}\pm0.30_{\rm sys}$	2.65	2.01	1.65
$0.73\pm0.03_{\rm stat}\pm0.11_{\rm sys}$	1.51	1.00	1.02
$0.63\pm0.03_{\rm stat}\pm0.09_{\rm sys}$	1.49	1.03	0.82
$0.84\pm0.03_{\rm stat}\pm0.12_{\rm sys}$	1.84	1.40	1.34
$0.51\pm0.02_{\rm stat}\pm0.07_{\rm sys}$	1.01	0.86	nc
$0.87\pm0.03_{\rm stat}\pm0.13_{\rm sys}$	2.69	1.55	1.73
$0.42\pm0.02_{\rm stat}\pm0.06_{\rm sys}$	1.54	0.98	1.33
$0.49\pm0.02_{\rm stat}\pm0.07_{\rm sys}$	1.60	1.02	1.15
	$\begin{array}{l} 1.54 \pm 0.06_{\rm stat} \pm 0.23_{\rm sys} \\ 1.21 \pm 0.05_{\rm stat} \pm 0.18_{\rm sys} \\ 2.05 \pm 0.08_{\rm stat} \pm 0.30_{\rm sys} \\ 0.73 \pm 0.03_{\rm stat} \pm 0.11_{\rm sys} \\ 0.63 \pm 0.03_{\rm stat} \pm 0.09_{\rm sys} \\ 0.84 \pm 0.03_{\rm stat} \pm 0.12_{\rm sys} \\ 0.51 \pm 0.02_{\rm stat} \pm 0.07_{\rm sys} \\ 0.87 \pm 0.03_{\rm stat} \pm 0.13_{\rm sys} \\ 0.42 \pm 0.02_{\rm stat} \pm 0.06_{\rm sys} \\ 0.49 \pm 0.02_{\rm stat} \pm 0.07_{\rm sys} \end{array}$	$\begin{array}{c c c c c c c c c c c c c c c c c c c $	$\begin{array}{c c c c c c c c c c c c c c c c c c c $





## A Better fit to the Gamma-Ray Sky



Fits are significantly improved, in particular in regions near the Galactic Center where there is significant kinematic gas information.



# Masking 1FIG Sources in the GC



Changing the point source catalog from the 3FGL to the 1FIG has only a negligible effect on the gamma-ray excess.

### The Effect on the Galactic center Excess (masking |b| < 2°)



Intriguingly, this persists even when the inner 2° are masked implying that analyses of small ROIs favors the excess.

### **The Low Energy Spectrum**



The Galactic Center models contain only a small preference for the convective winds, and the spectrum and intensity of the Galactic center excess component remains resilient.

### **Convection in the Galactic Center**



This increases the best fit value of  $f_{\rm H2}$  for the GC data, bringing this value into agreement with the global best fit value.

Models with a GCE component still prefer slightly lower values of  $f_{\rm H2}$ , but these have increased to 0.2 as well.

#### The Galactic Center Excess Morphology (masking |b| < 2°)



The deviations from typical NFW profiles are more extreme when the |b| < 2° is masked from the analysis, with a shallower emission profile preferred by the data.

## Morphology in the Galactic Center



GC

For the Galactic Center analysis, the morphology of the excess component remains relatively robust

# Analysis Far from the GC



Analysis regions far from the GC also show an excess – not much star formation occurs a few degrees above the Galactic plane.

Calore et al. (2014, 1409.0042)



### Comparison to Cygnus-X





Unprocessed map of 1.0 to 3.16 GeV gamma rays

Known sources removed

# **A Fermi-LAT Based Example**



Total Gamma-Ray Flux (>1 GeV) in inner 1° is 1.1 x 10<sup>-9</sup> erg cm<sup>2</sup> s<sup>-1</sup>

Approximately half of this emission is produced along the line of sight towards the GC, and thus we approximate the total gamma-ray luminosity of the central one degree to be  $5 \times 10^{36}$  erg s<sup>-1</sup>

### Reacceleration

A 100 M<sub>o</sub> molecular cloud with a gas density of 200 cm<sup>-3</sup> in a 10  $\mu$ G B-field supports a Kinetic Energy transfer (via magnetic turbulence) of 5.0 x 10<sup>32</sup> erg s<sup>-1</sup>.

Given a total molecular cloud mass of  $10^7 \text{ M}_{o}$ , the total reacceleration of cosmic-ray electrons is 5.0 x  $10^{37}$  erg s<sup>-1</sup>.

If these electrons are trapped until losing their energy to bremsstrahlung and synchrotron emission with a standard spectrum):
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6.1 x 10<sup>36</sup> erg s<sup>-1</sup>

# A Fermi Bubbles Component?

When the excess floats to the best fit morphological configuration, much of the excess intensity returns.

Most importantly, the over subtraction issue at low energies is fixed.



## The Excess is Degenerate with $f_{\rm H2}$



Models with no dark matter universally prefer  $f_{H2} \sim 0.2$  for the 40°x40° region surrounding the GC.

Models with an NFW emission template prefer  $f_{H2} \sim 0.1$ .

The reduction in the normalization of the NFW template is ~1.5 for  $f_{H2} \sim 0.1$ , instead of a factor of 3 at  $f_{H2} \sim 0.2$ .

# Dark Matter Annihilation?



$$\left(\frac{\Omega_{\chi}}{0.2}\right) \simeq \frac{x_{\text{f.o.}}}{20} \left(\frac{10^{-8} \text{ GeV}^{-2}}{\sigma}\right)$$

$$\langle \sigma v \rangle \sim 10^{-8} \text{ GeV}^{-2} \left( 3 \times 10^{-28} \text{ GeV}^2 \text{ cm}^2 \right) \ 10^{10} \ \frac{\text{cm}}{\text{s}} = 3 \times 10^{-26} \ \frac{\text{cm}}{\text{s}}$$



A particle with a weak interaction cross-section and a mass on the weak scale is expected to naturally obtain the correct relic abundance through thermal freeze-out in the Early Universe.