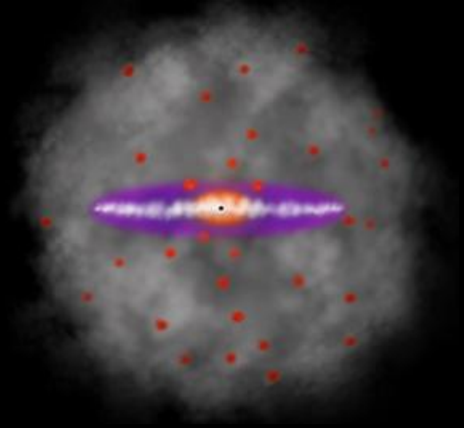
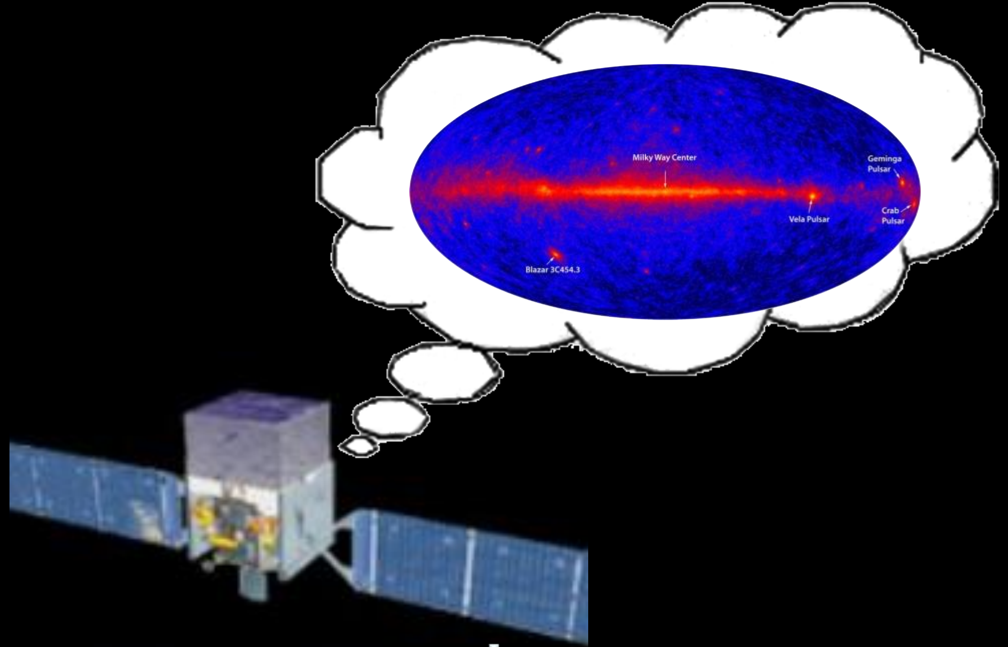




Searching for Dark Matter in the Fermi Era

Tim Linden[1]

with Stefano Profumo[1]

Brandon Anderson[1] and Jennifer Siegal-Gaskins[2]

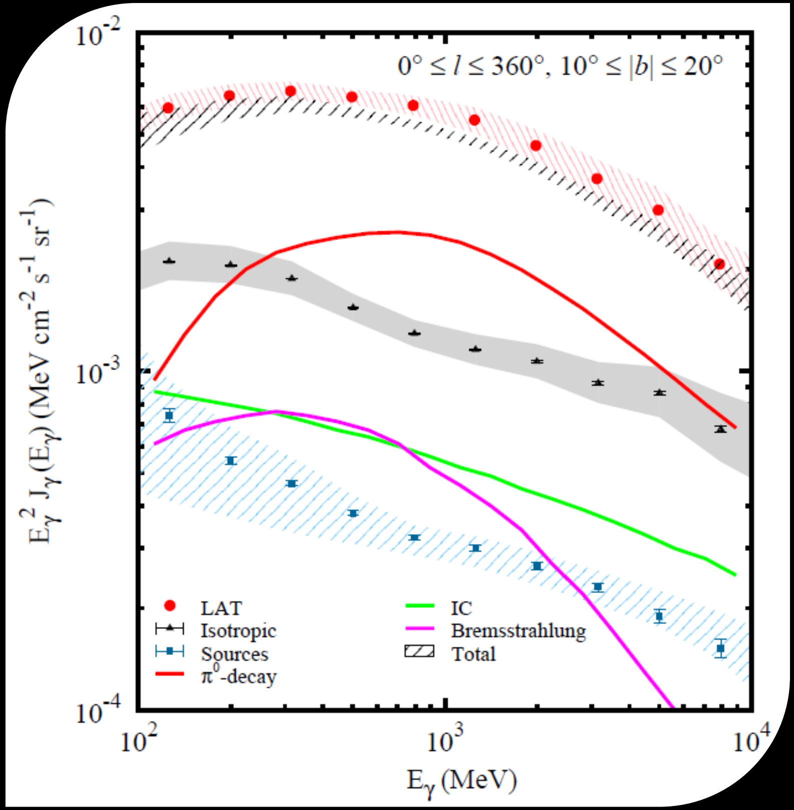
1 University of California – Santa Cruz
2 The Ohio State University

Cosmo Club UC – Santa Cruz April 21, 2010

Outline

- Are we already seeing dark matter?
- Difficulties in extracting a dark matter signal
- How could we determine whether an unknown signal is due to dark matter annihilation?

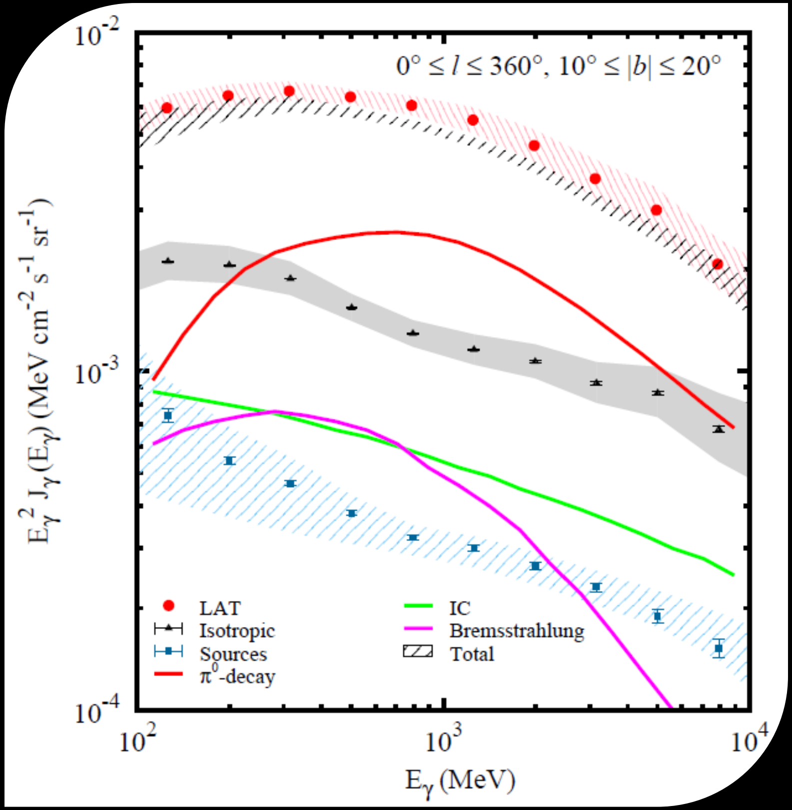
Observations of Diffuse Emission



T. Porter, 2009 (0907.0294)

- Many sources:
  - Extragalactic
  - π^0 decay
  - Inverse Compton Scattering
  - Bremsstrahlung Emission
  - Unresolved Point Sources

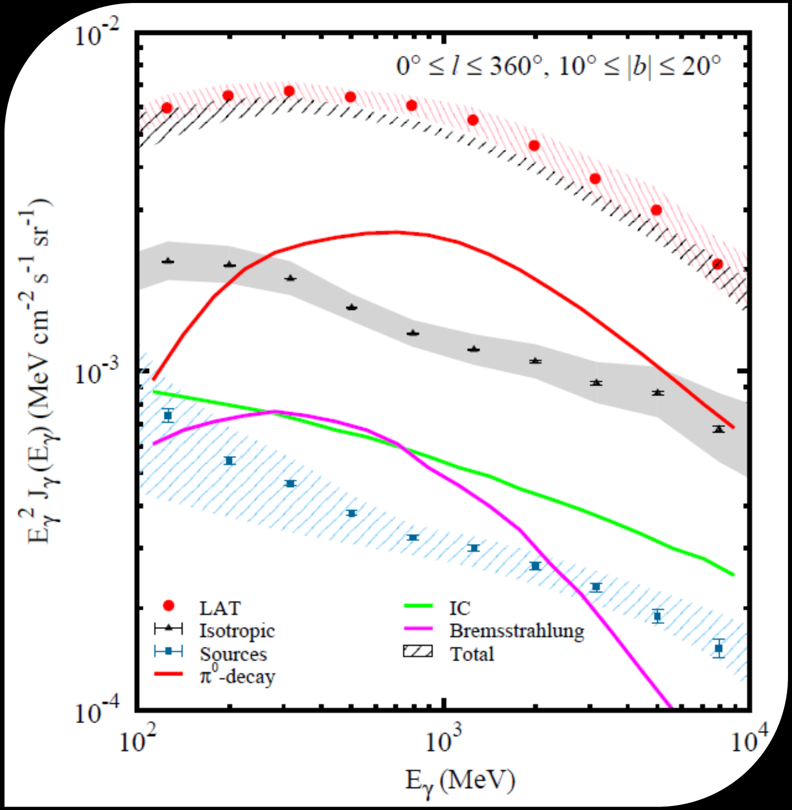
Observations of Diffuse Emission



T. Porter, 2009 (0907.0294)

- Many sources:
  - Extragalactic
    - Isotropic Source Classes
    - Sources include blazars, starburst galaxies
  - π^0 decay
  - Inverse Compton Scattering
  - Bremsstrahlung Emission
  - Unresolved Point Sources

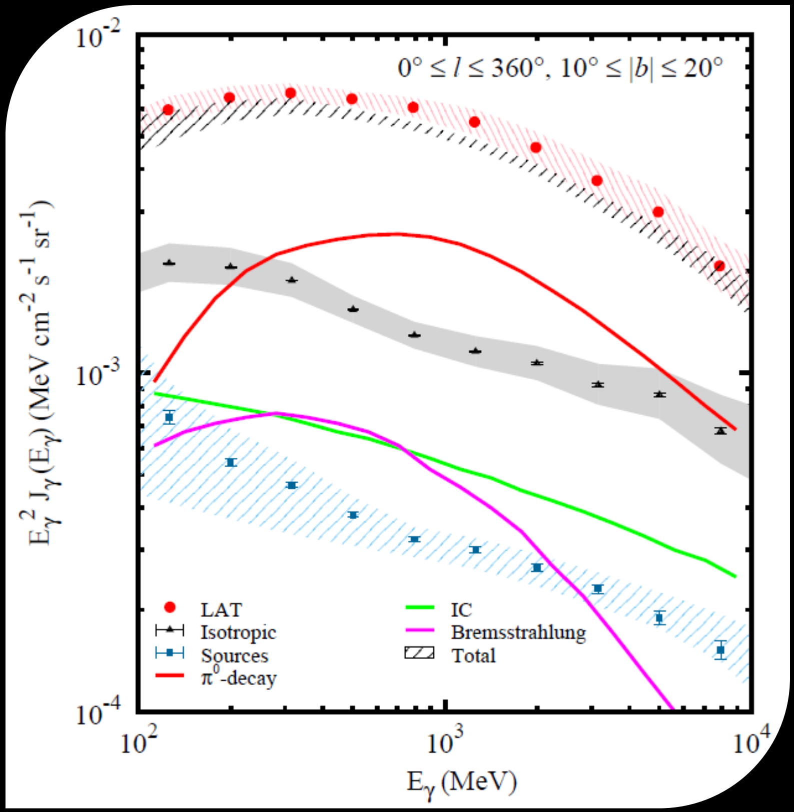
Observations of Diffuse Emission



T. Porter, 2009 (0907.0294)

- Many sources:
  - Extragalactic
  - π^0 decay
    - Proton collisions with galactic dust
    - Well defined emission spectra
  - Inverse Compton Scattering
  - Bremsstrahlung Emission
  - Unresolved Point Sources

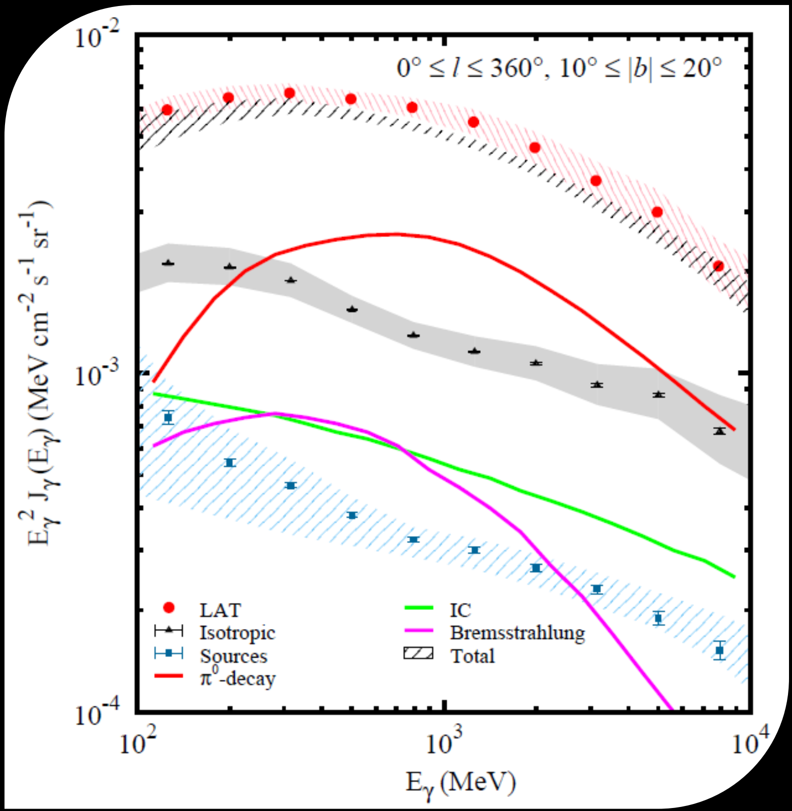
Observations of Diffuse Emission



T. Porter, 2009 (0907.0294)

- Many sources:
  - Extragalactic
  - π^0 decay
  - Inverse Compton Scattering
    - Interactions of charged leptons with Interstellar radiation field
  - Bremsstrahlung Emission
  - Unresolved Point Sources

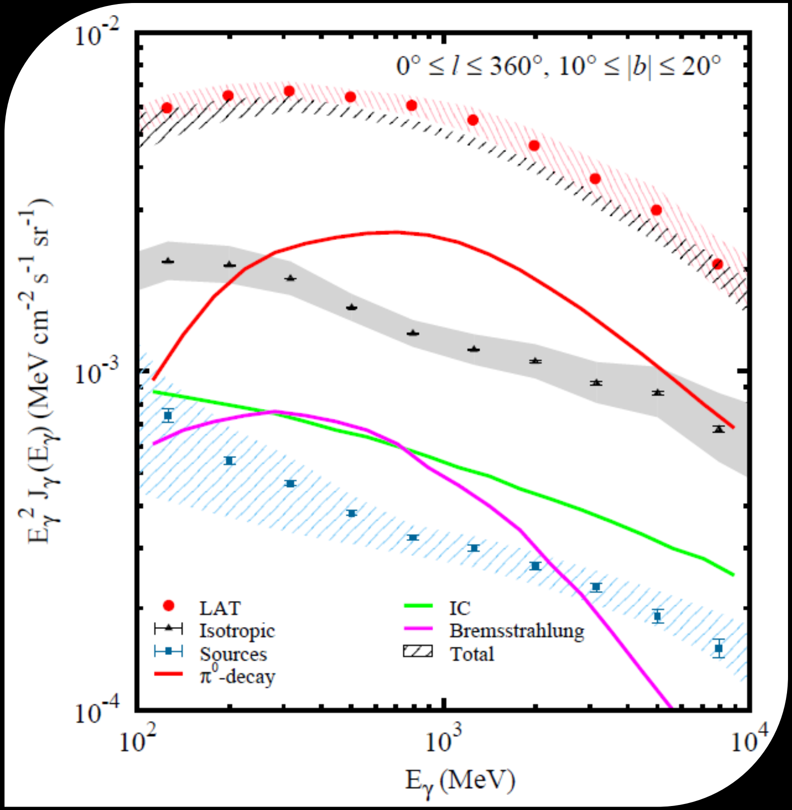
Observations of Diffuse Emission



T. Porter, 2009 (0907.0294)

- Many sources:
  - Extragalactic
  - π^0 decay
  - Inverse Compton Scattering
  - Bremsstrahlung Emission
    - Relatively weak sources at high energies and away from the galactic center
  - Unresolved Point Sources

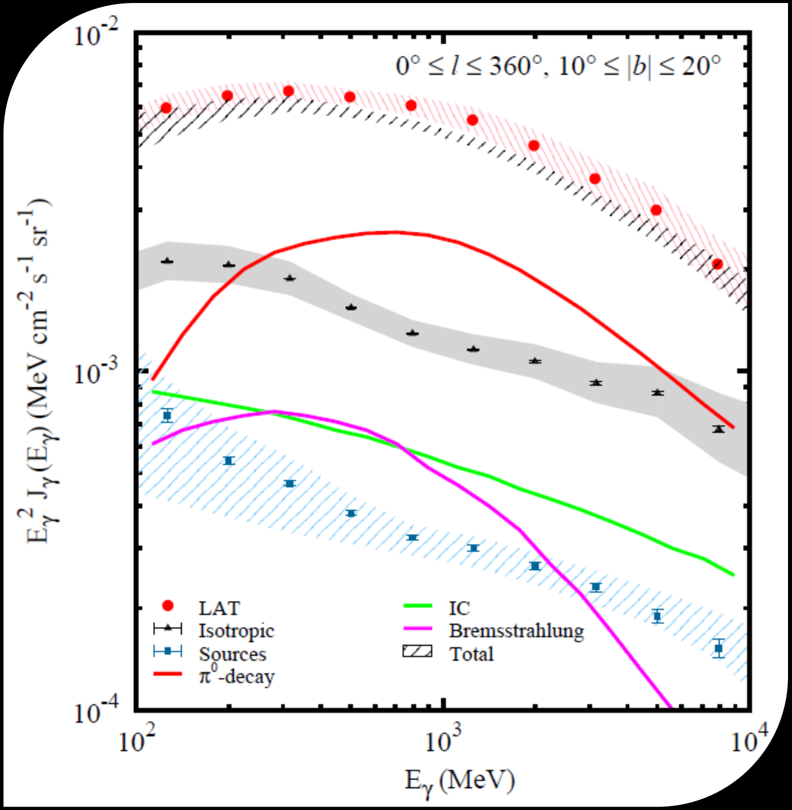
Observations of Diffuse Emission



T. Porter, 2009 (0907.0294)

- Many sources:
  - Extragalactic
  - π^0 decay
  - Inverse Compton Scattering
  - Bremsstrahlung Emission
  - Unresolved Point Sources
    - Systematic error – Intensity and spectra changes over time

Is There Any Room Left for DM?

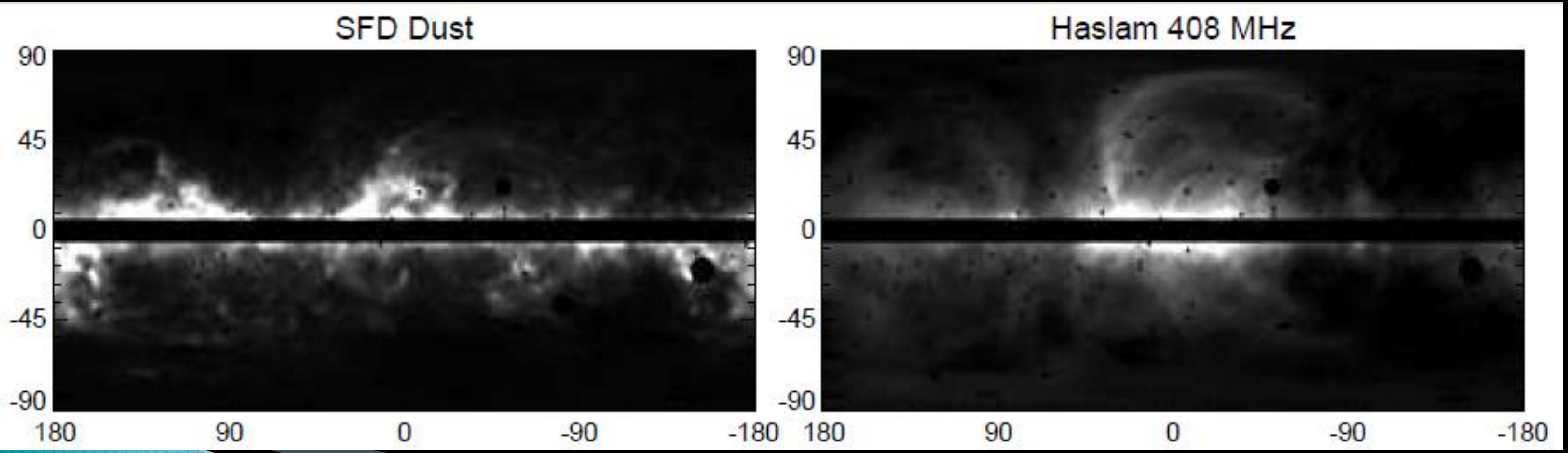


T. Porter, 2009 (0907.0294)

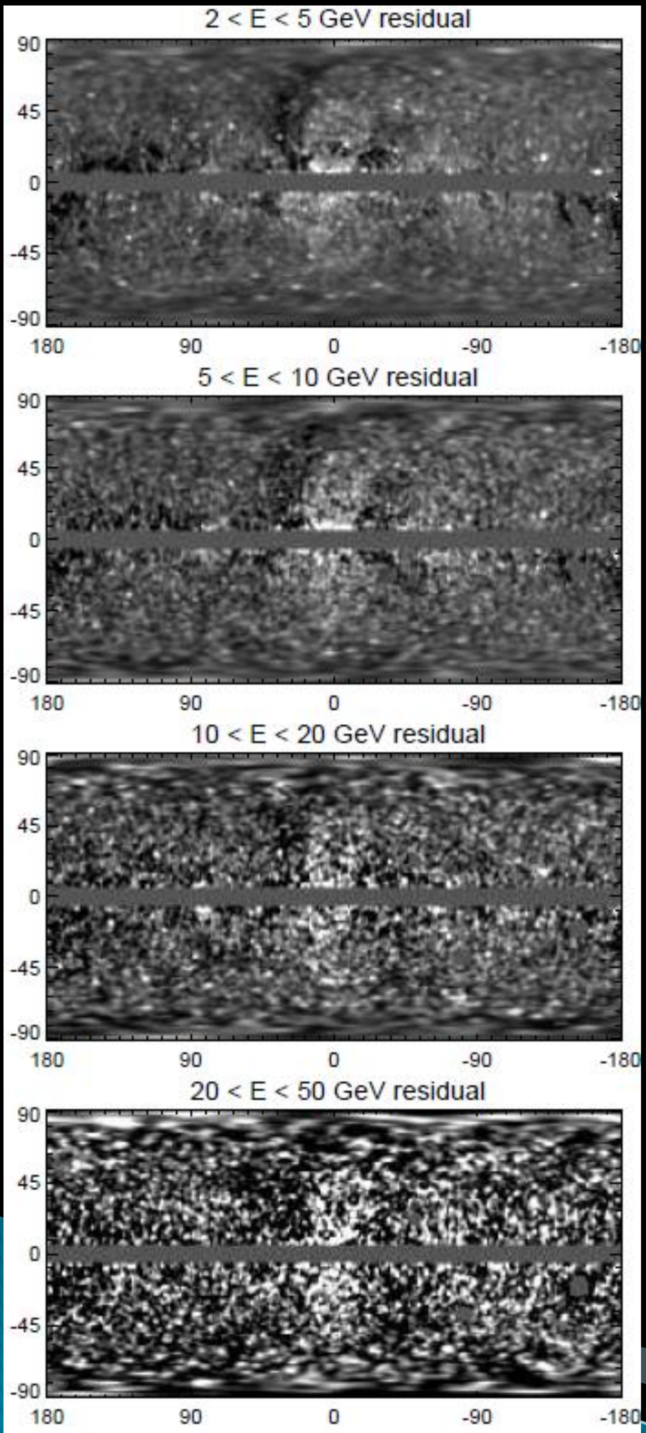
- How large are the astrophysical uncertainties in each of these background signals?
- Playing a very different game than direct dark matter detection (e.g. CDMS)
  - Should see millions of events
  - But no background rejection

Modeling Astrophysical Sources

- Dobler et al. (2009) created models for the morphology of these astrophysical components
  - Point sources subtracted from 3-month Fermi catalog
  - SFD Dust Map for π^0 decay
  - Haslam 408 Mhz map for ICS
  - Residual Map mean subtracted to eliminate isotropic



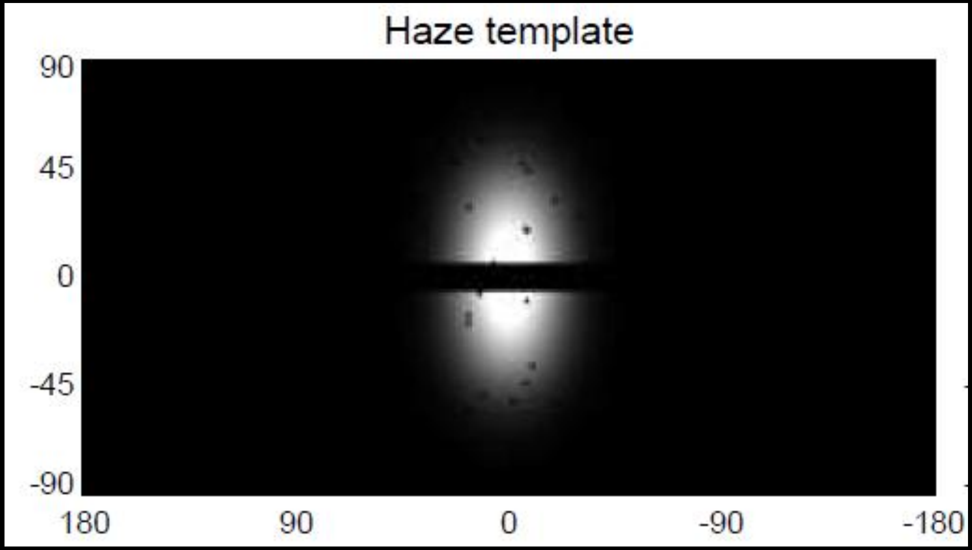
Dobler et al (2009) (0910.4583)



An Enticing Residual!

- Dobler et al. finds a significant residual when these maps are applied
- Residual has a pronounced morphology above and below the galactic center

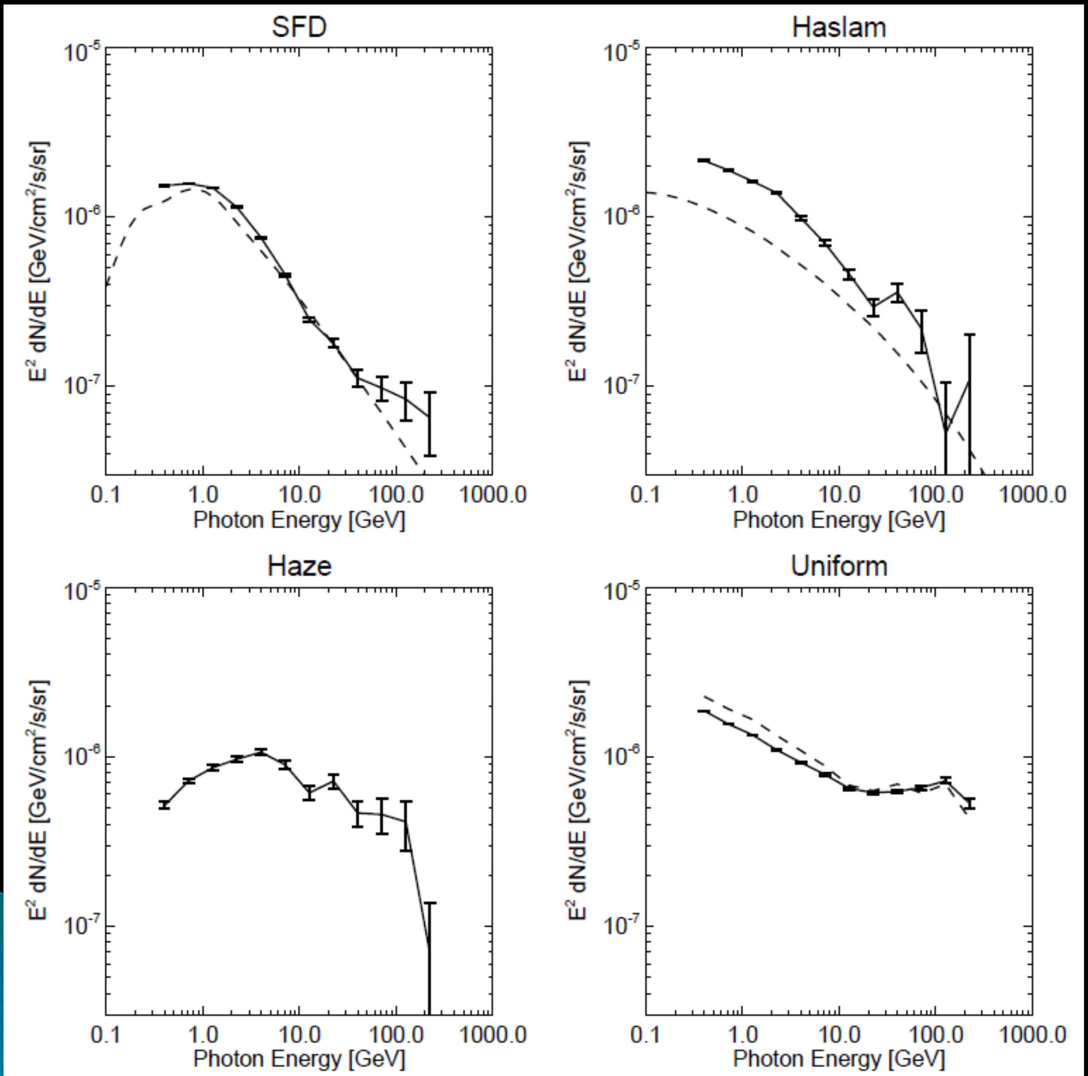
Modeling this residual



- To map this residual, Dobler et al add an ad hoc template

This template is a bivariate gaussian with latitudinal scale height $\sigma_b = 25^o$, and a longitudinal scale height of $\sigma_l = 15^o$

These residuals are large!

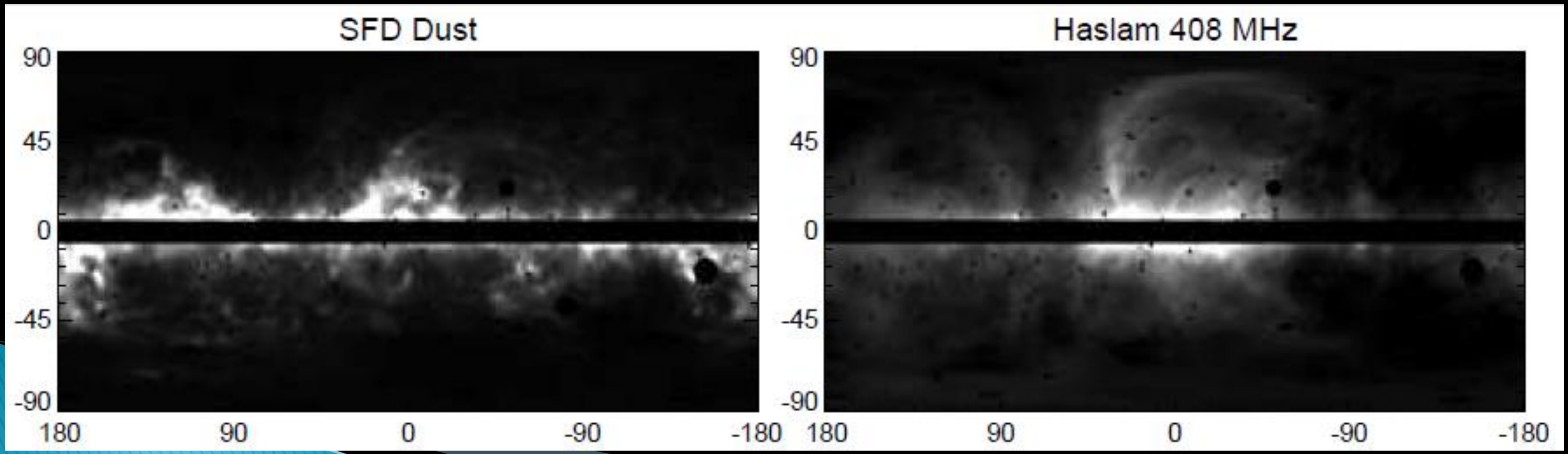


- The haze template is **co-dominant** with the isotropic background above 10 GeV
- Other templates have reasonable spectra

$|l| < 15^o \quad -30^o < b < -10^o$

But do we trust these templates?

- SFD Dust Map for π^0 decay
  - Dust is a reasonable tracer for galactic gas
  - Gas acts as the target of energetic protons

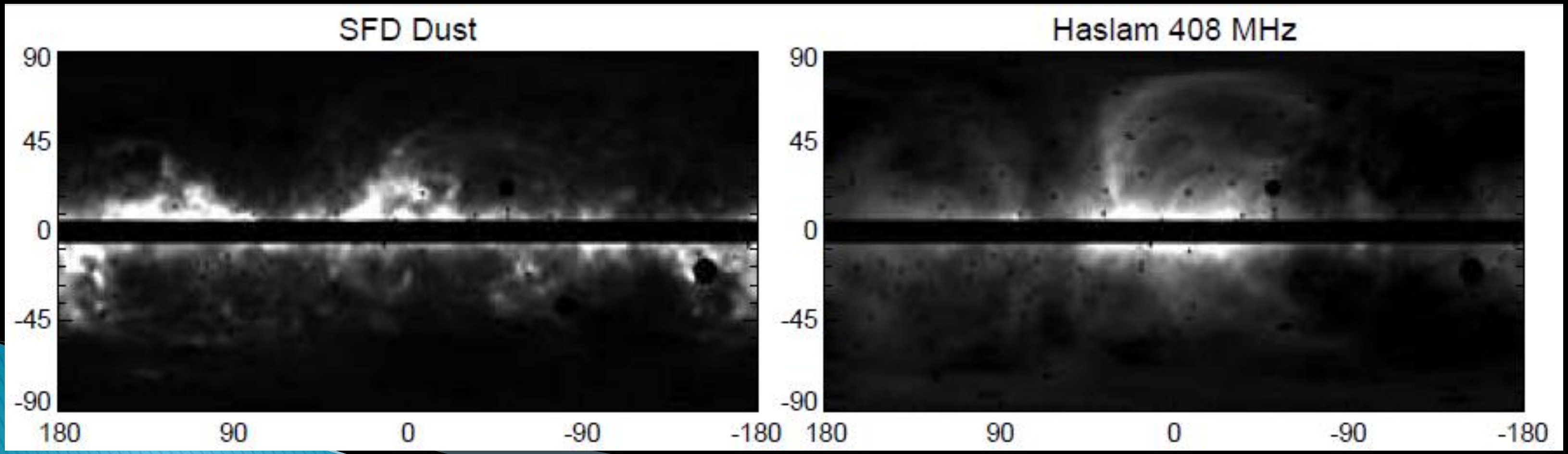


But do we trust these templates?

- SFD Dust Map for π^0 decay
  - But the cosmic ray distribution is not isotropic!

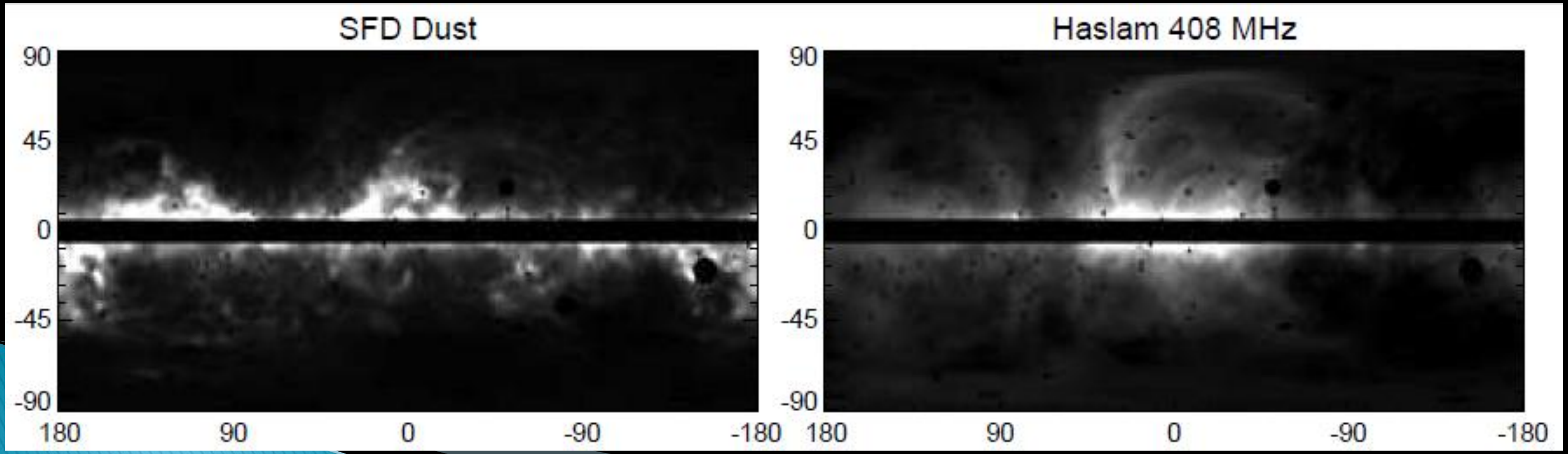
$$\Phi_{\pi 0} = \text{BEAM x TARGET}$$

Cosmic rays → BEAM; Galactic Gas → TARGET



But do we trust these models?

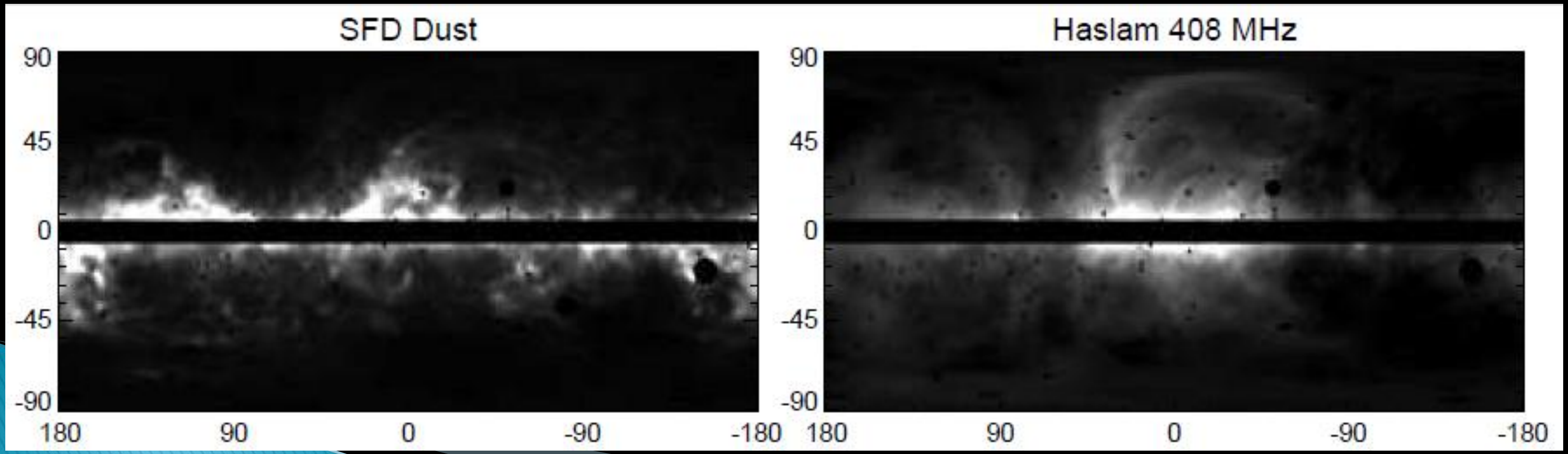
- Haslam 408 Mhz skymap for inverse Compton scattering
  - At 408 Mhz, the radio sky should be dominated by synchrotron of energetic leptons.
  - These same leptons should create γ-ray emission due to ICS of the interstellar radiation field.



But do we trust these models?

- Haslam 408 Mhz skymap for inverse Compton scattering
  - The morphology of the interstellar radiation field is <u>not</u> the same as the morphology of the galactic magnetic fields.

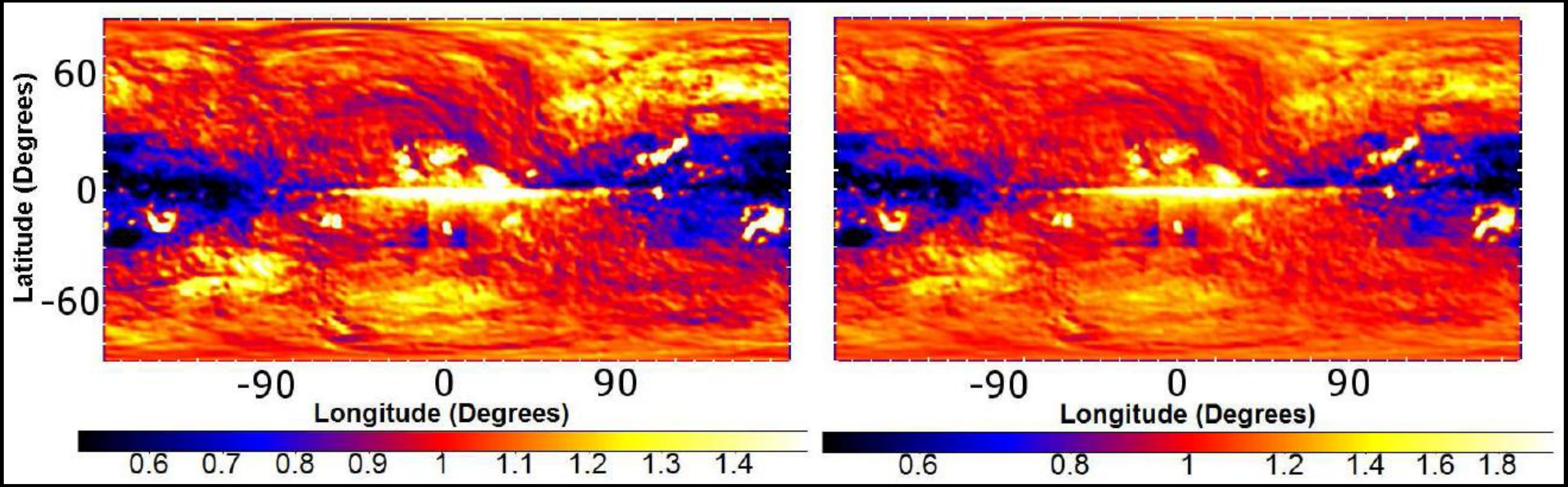
$$\frac{\text{ISRF}}{B^2} \neq 1$$



Our Setup

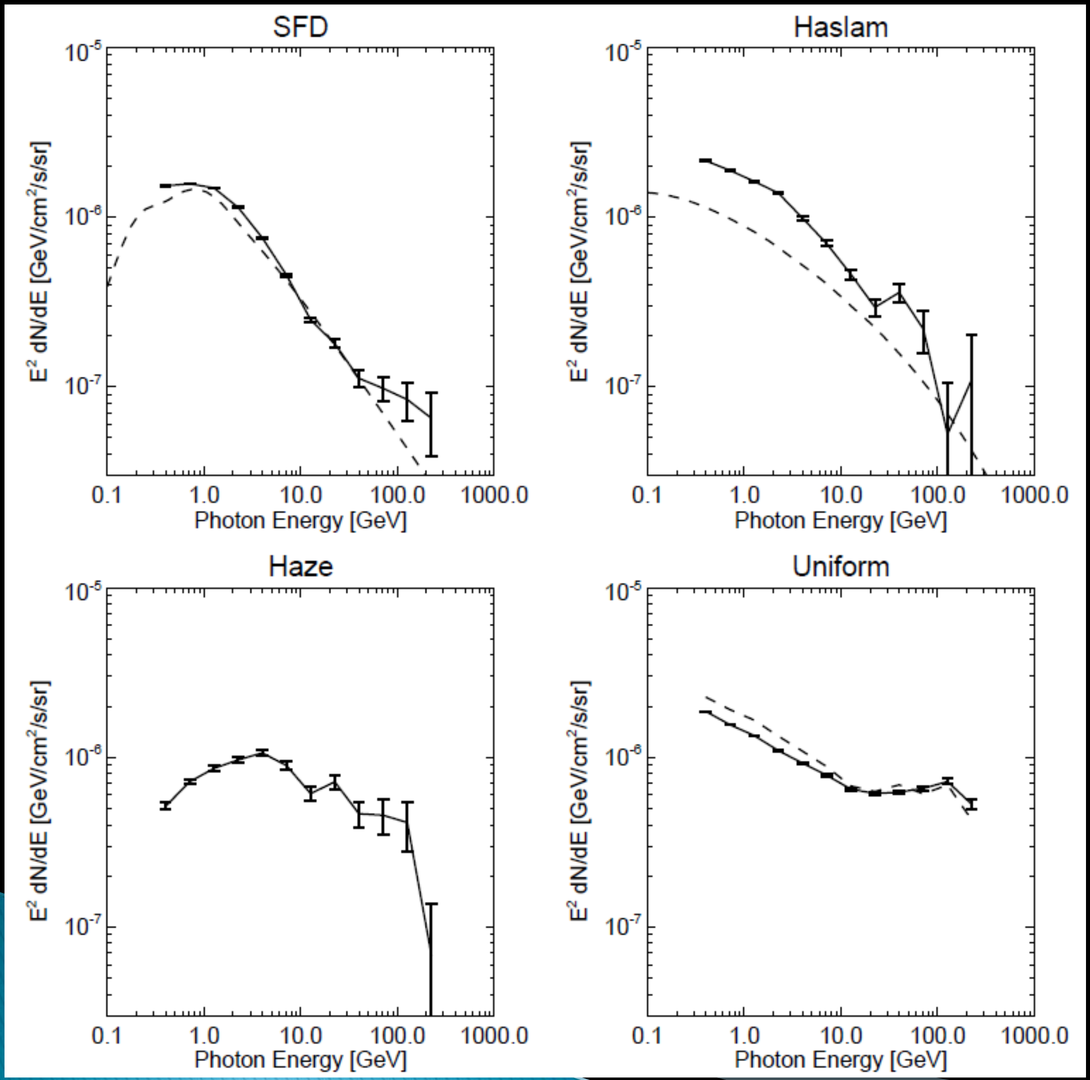
- We use GALPROP models to test the morphological consistency of:
  - 1.) The π^0 decay morphology and the input gas map
  - 2.) The ICS decay morphology and the 408 Mhz synchrotron morphology

π^0 Decay divided by Gas Map



- The resulting skymap has a haze-like morphology that can be fit with a Gaussian 17% as strong as the estimated π^0 skymap

Important at low energies



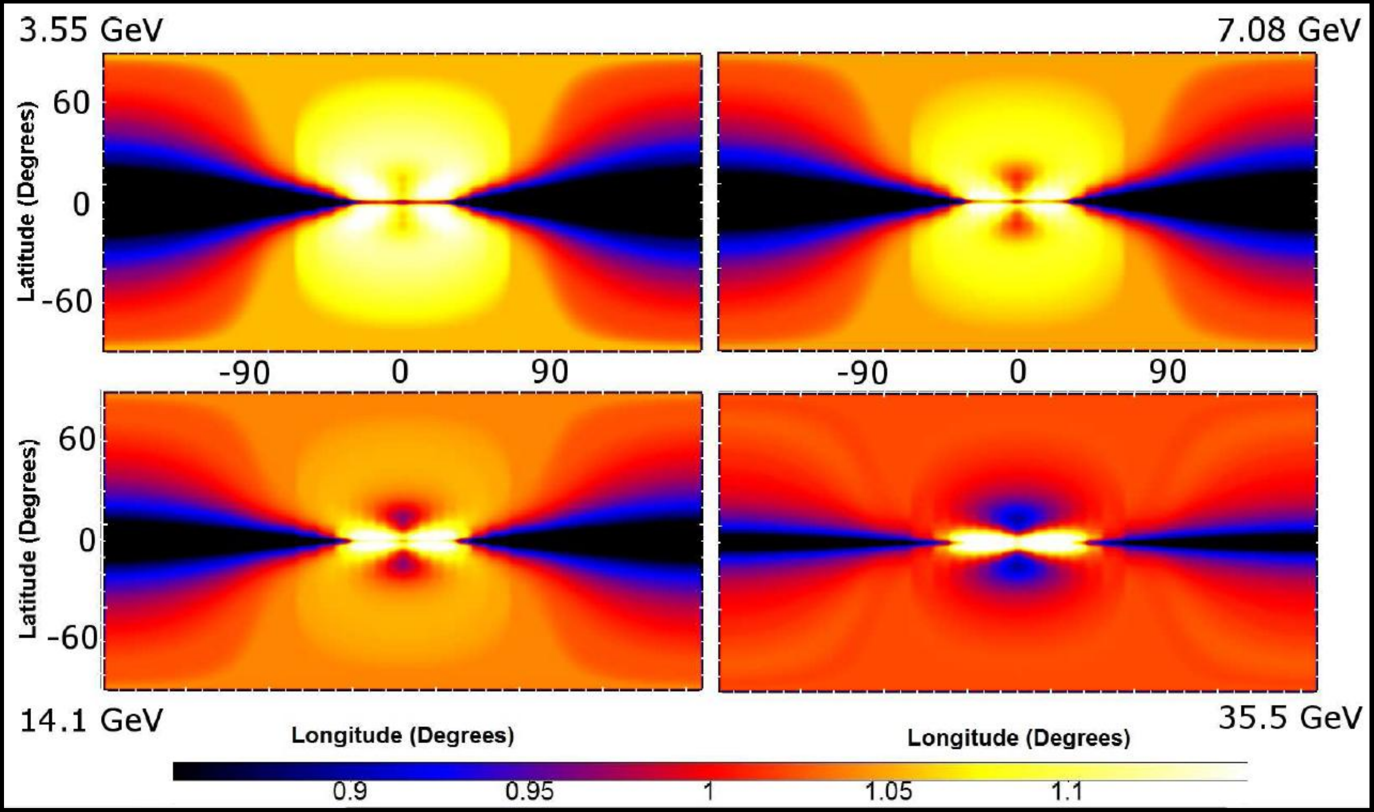
- The π^0 decay morphology is dominant up to several GeV
- A 17% Gaussian residual can explain the majority of the haze around 1 GeV

ICS divided by Synchrotron

- The ratio of the ICS map to the synchrotron map also show a haze morphology

Highly energy dependent

Low intensity 16% to -4%



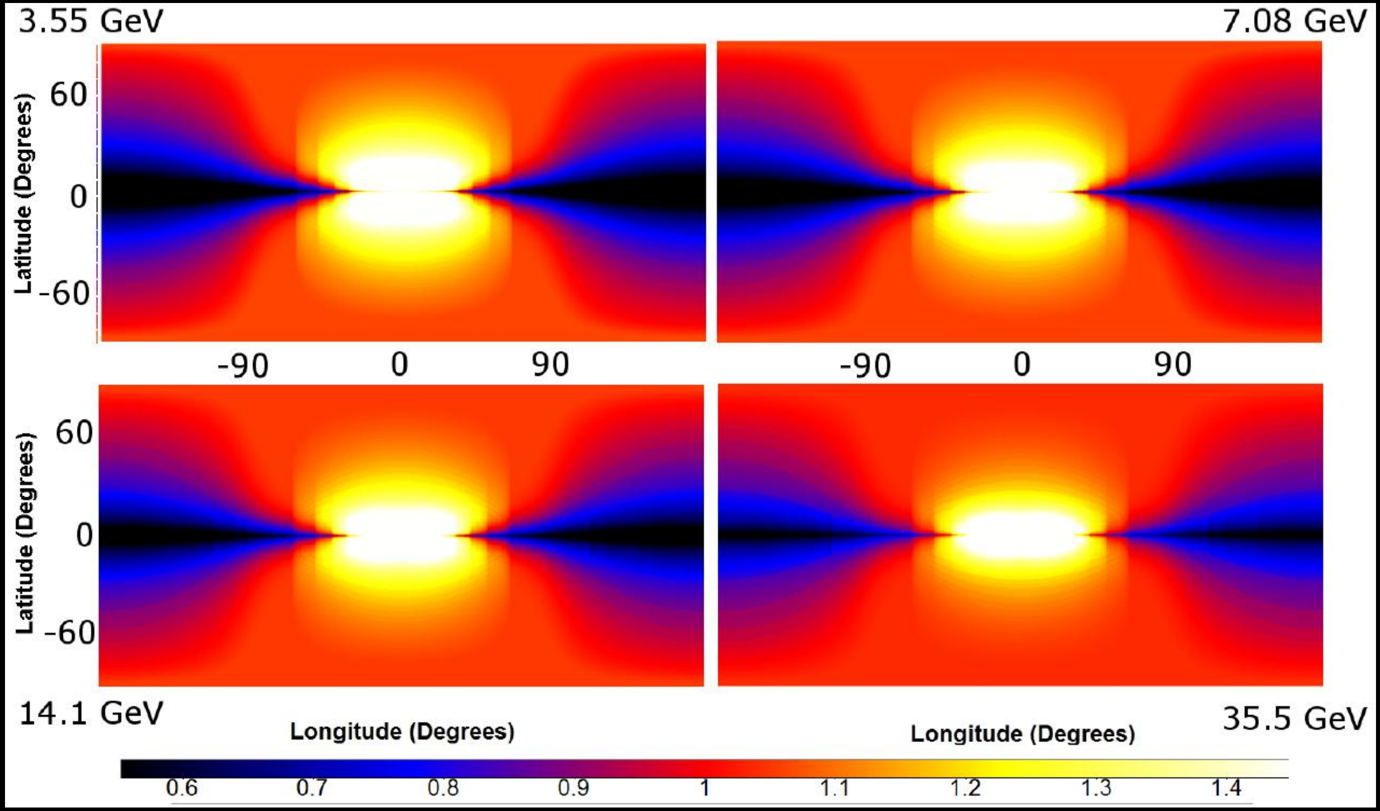
$B = 11.6\ \mu G\ e^{-r/10\ kpc\ -z/2kpc}$

ICS / Sync is very uncertain

- By slightly altering the magnetic field morphology, we can create larger deviations

Same energy dependence as before

Large Gaussian errors (54% to 40%)



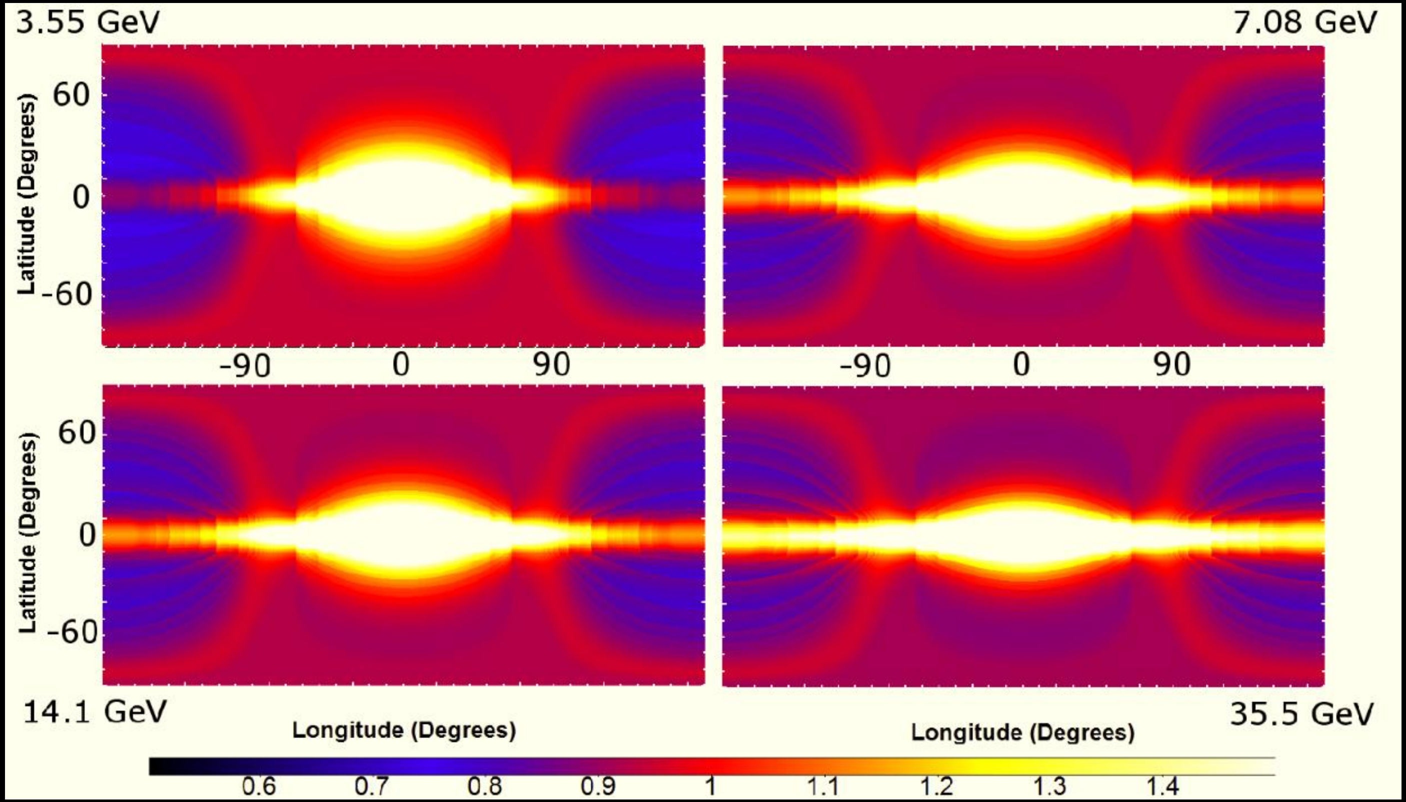
$B = 11.6\ \mu G\ e^{-r/20\ kpc\ -z/2kpc}$

ICS / Sync is very uncertain

- By slightly altering the magnetic field morphology, we can create larger deviations

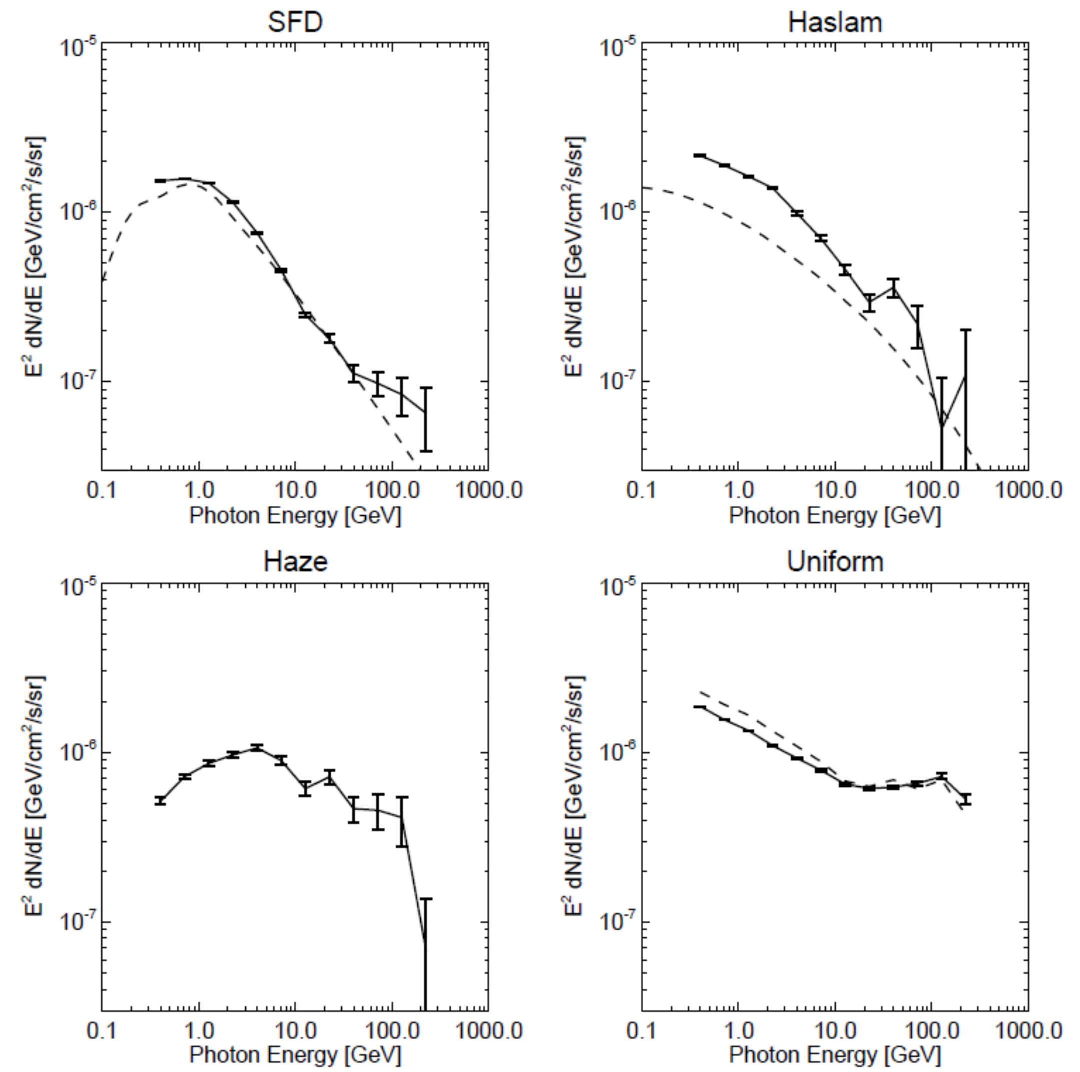
Flat magnetic field model

This model was used to generate the WMAP haze morphology



$B = 10.0\ \mu G\ e^{-r/99.9\ kpc\ -z/99.9kpc}$

Also Important at low energies



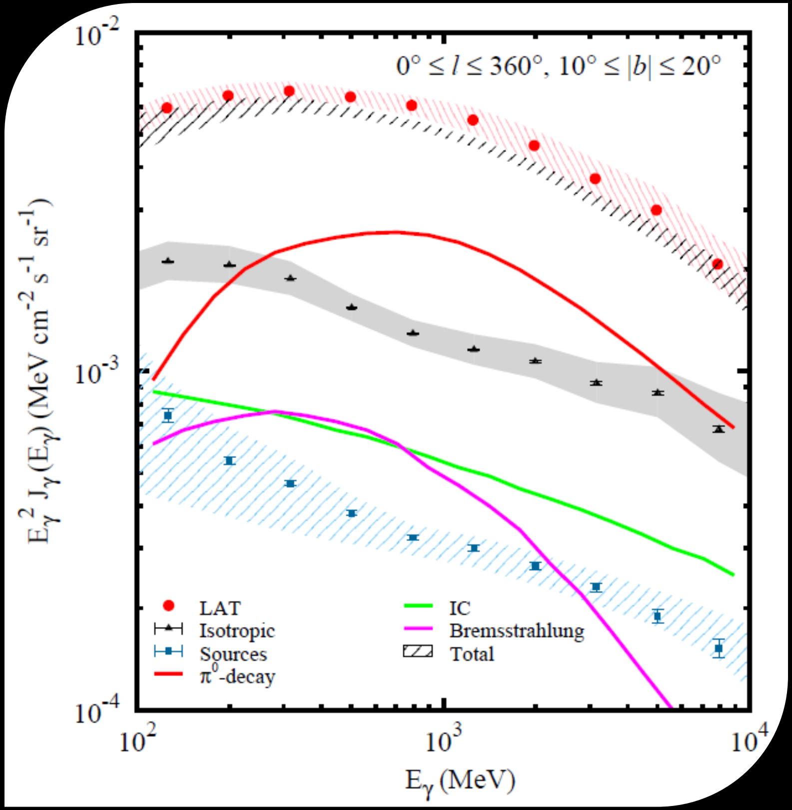
- Expected ICS signal is a significant fraction of the haze residual below 10 GeV
- This appears to leave a discrepancy at high energies, but eliminates it below ~10 GeV

First Conclusions

- The current analysis of the Fermi-Haze is insufficient to determine either the intensity or the spectrum of any Fermi residual

- Early measurements suggest that the Fermi haze at low energies (<10 GeV) could be entirely explained with theoretically correct templates for pi0 and astrophysical ICS emission.

A More Direct Analysis

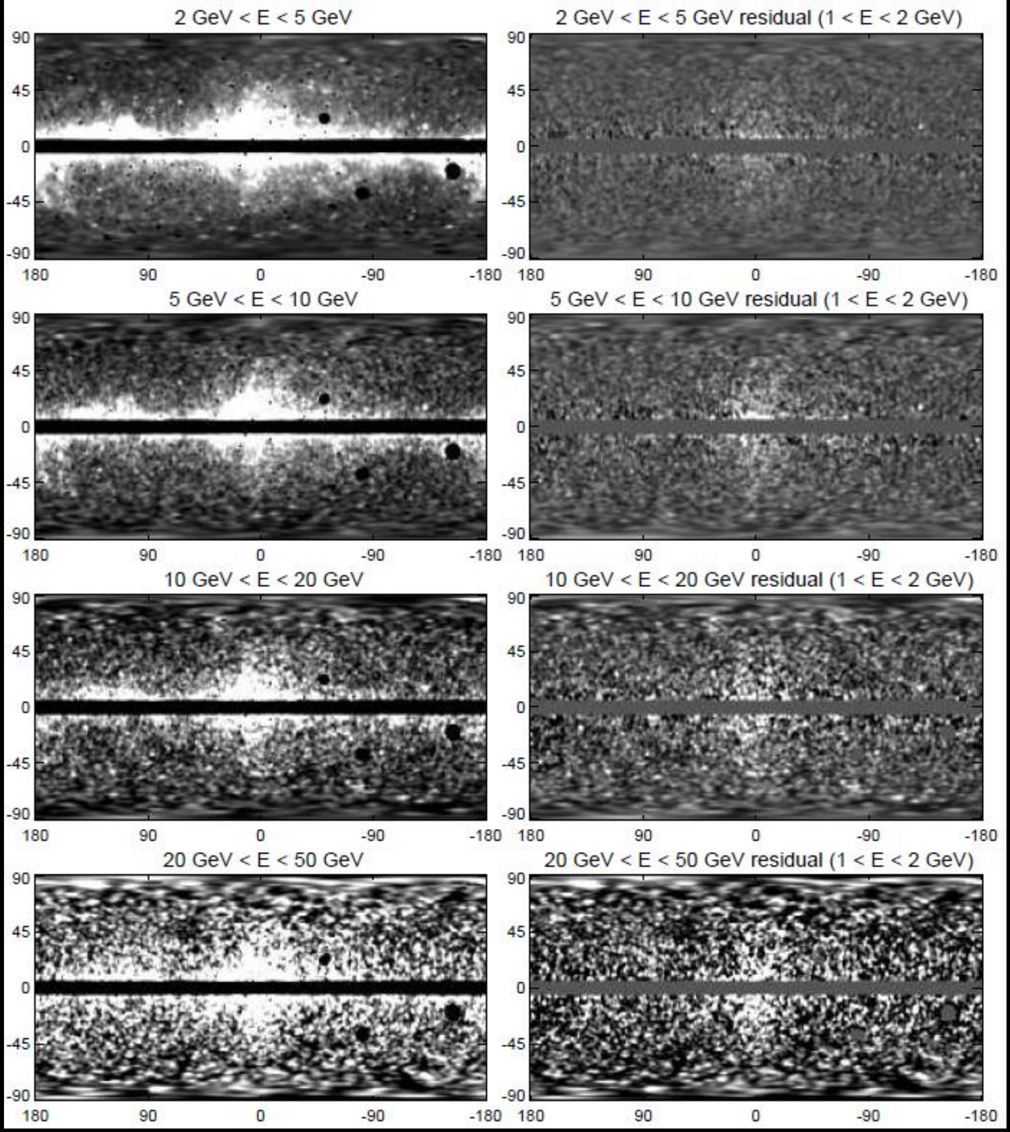


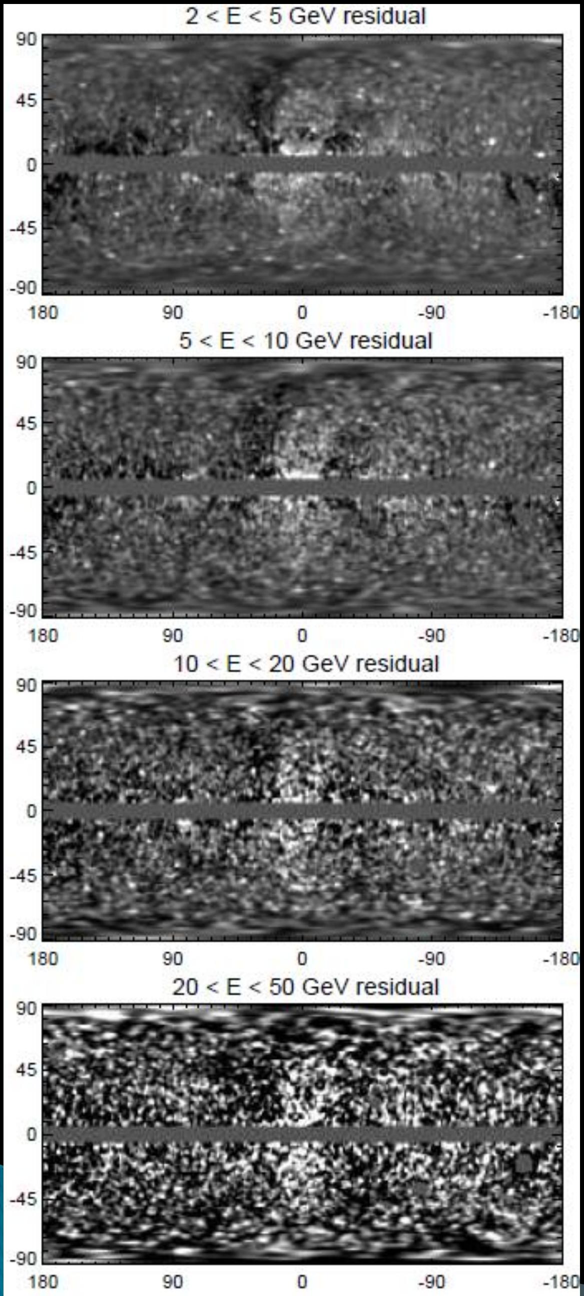
T. Porter, 2009 (0907.0294)

- At 1-2 GeV, emission from π^0 decay should be highly dominant
- Note that the π^0 decay morphology should be constant as a function of energy

Residuals!

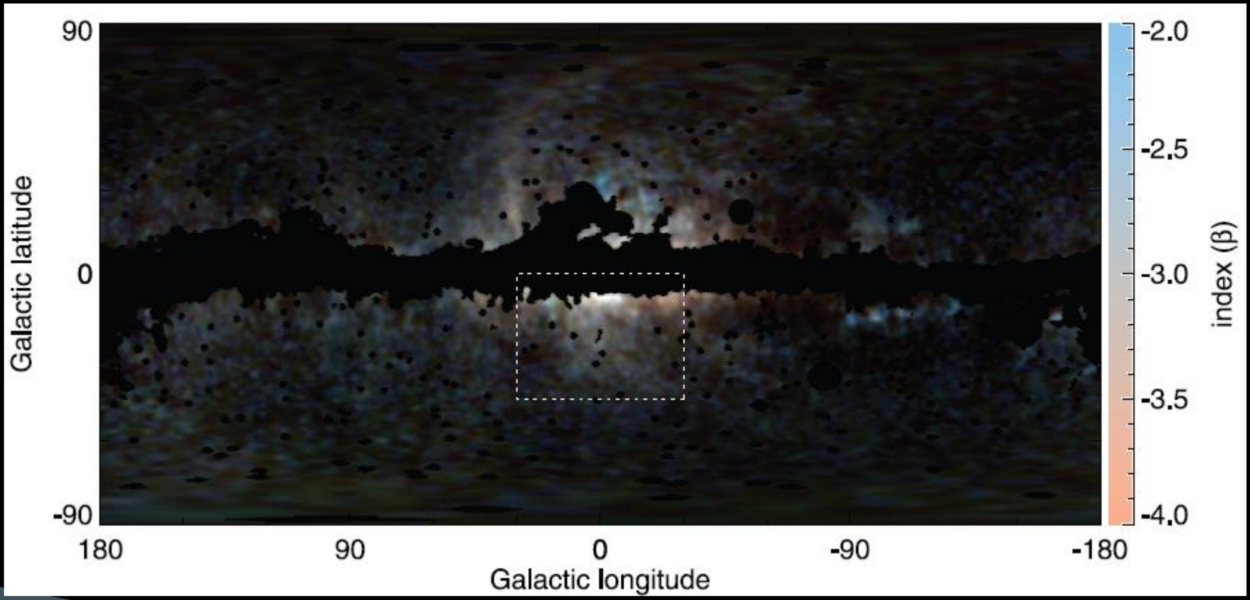
- Process:
  - 1.) Subtract out the morphology of the 1-2 GeV map
  - 2.) Find the morphology of the residual
- Dobler et al. find there is still a visible haze
- Some/Most of this haze is likely astrophysical ICS
- Can't determine an intensity/spectrum





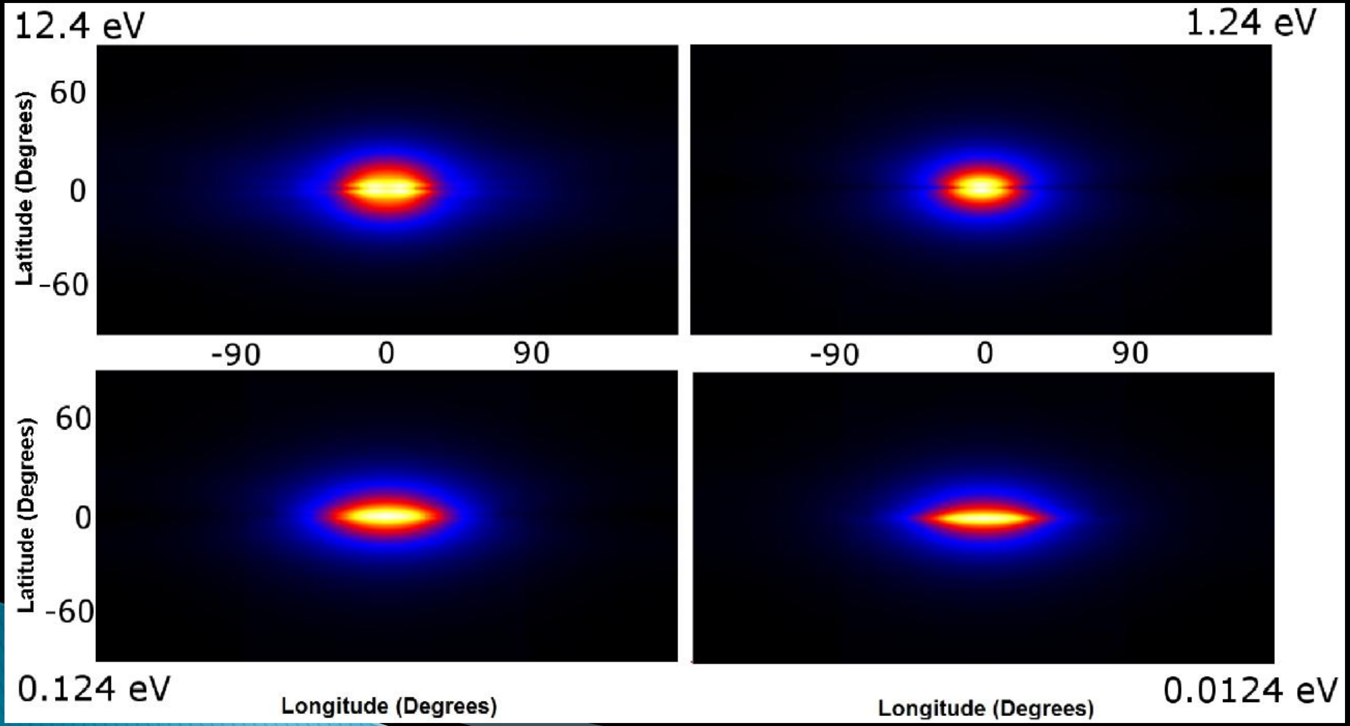
Understanding Residuals

- Dobler et al note a correlation between the Fermi haze and the WMAP haze
- Find that this suggests a new primary electron source near the galactic center



Understanding Residuals

But it is difficult to model the Fermi haze with only a new lepton input class, as the interstellar radiation function falls off too quickly.

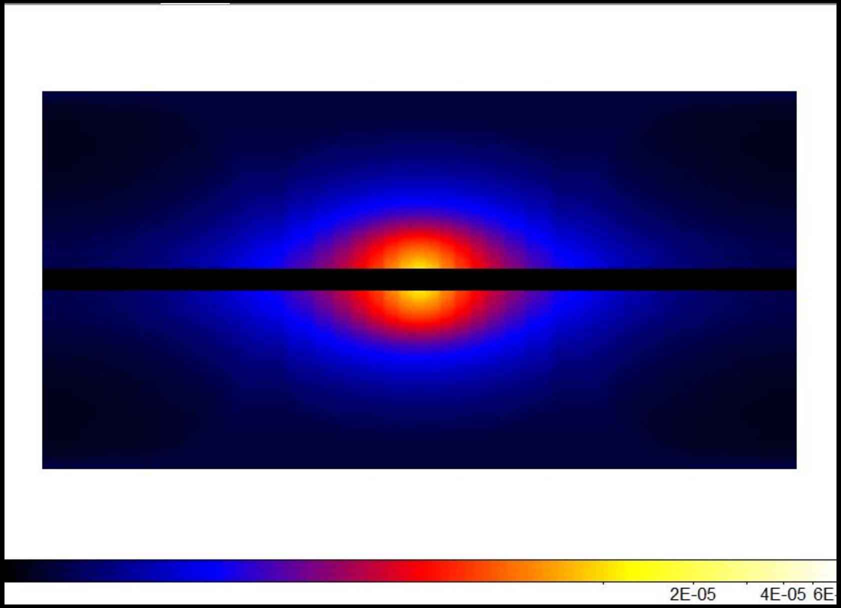


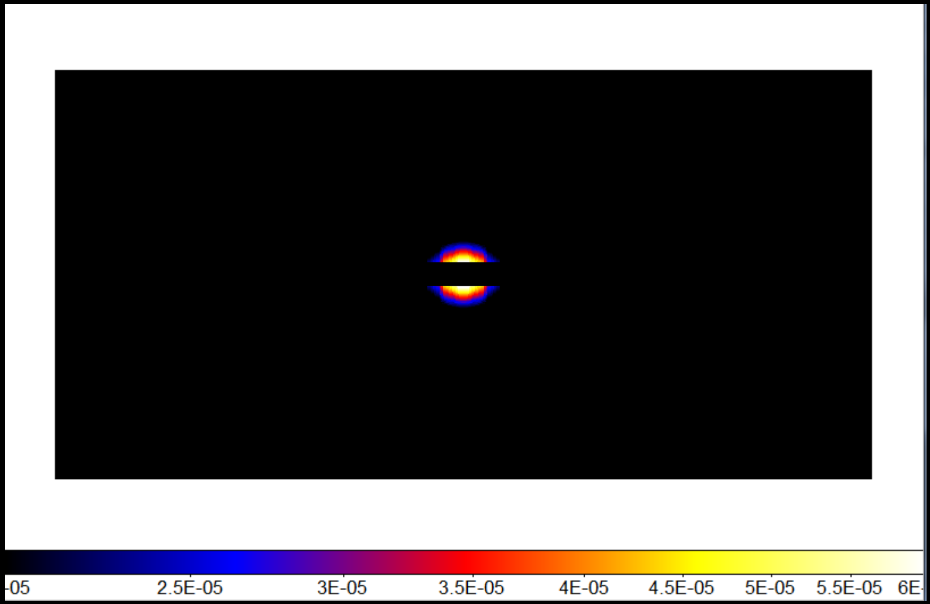
Between 10º to 30º latitude, the ISRF dims by between 63-72%

The longitudinal to latitudinal extent moves from 5-4 to 9-1 at low energies

DM Interpretation?

Simulated Dark matter models produce a haze that decays much too quickly at high latitudes to match the observed haze.





1500 GeV -> $\mu^+\mu^-$
Snapshot at 8 GeV

Possible Sources

- There are quite a few possible sources for this residual:
  - Nearby sources
  - Jets from galactic center
  - Magnetic anomalies (e.g. Loop 1 – Casandjian et al. 2009, 0912.3478)
  - Changes in the Interstellar radiation field
  - Energy dependence changes in diffusion parameters
- But it is very difficult to match emission using purely diffuse sources or spherical distributions, as they will always have more longitudinal extent.

How do we distinguish?

- New methods are necessary for distinguishing between various emission mechanisms
    - Multi-wavelength studies
    - Anisotropic studies

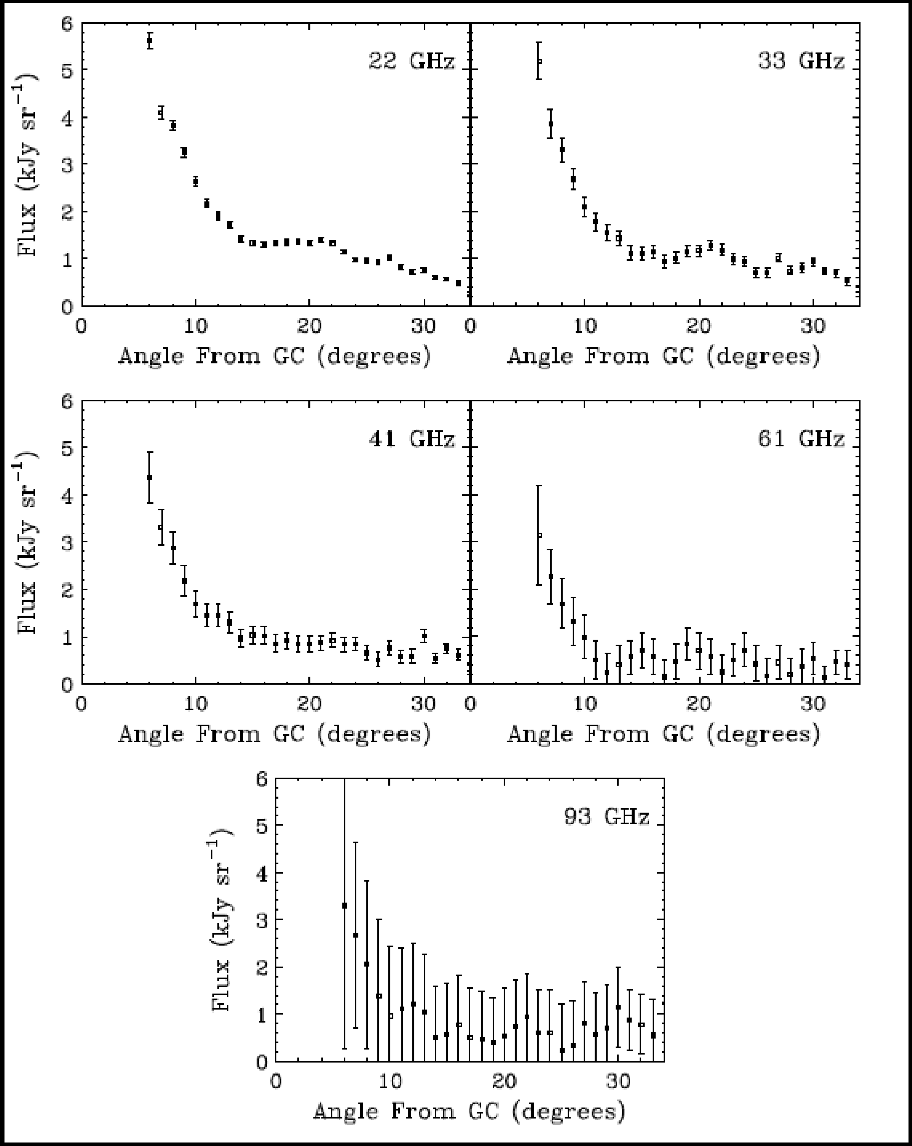
Multi-wavelength Studies

- Several Important cross-checks:
  - WMAP – Energetic leptons should also produce microwave radiation through synchrotron in the galactic B-field
  - ROSAT – SNR and galactic anomalies may produce X-ray signatures
  - PAMELA – New primary sources may match the observed positron/electron spectrum
  - HESS – Very high energy γ-ray's should match these observations

The WMAP Haze

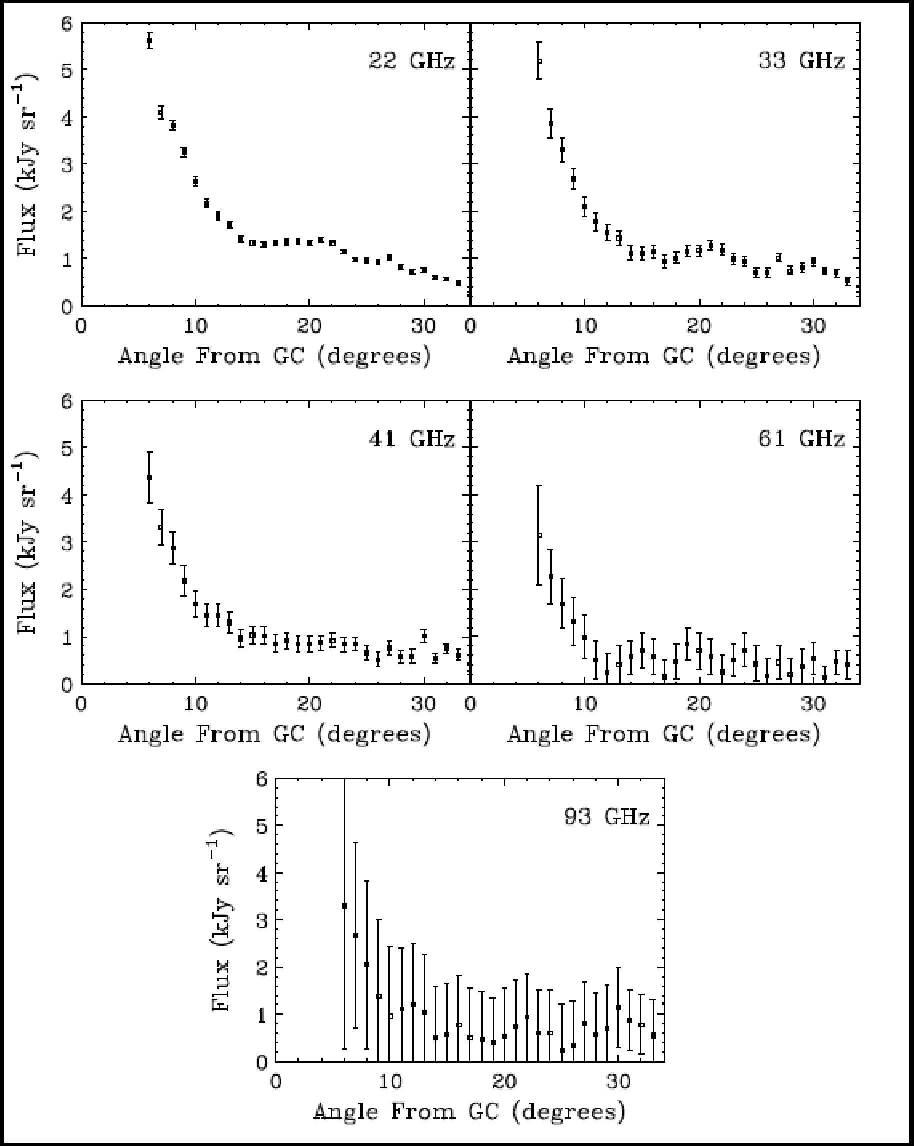
- Finkbeiner (2004) pointed out an unexplained residual in the WMAP dataset

- The existence of this residual is controversial, and is not detected by the WMAP team (Gold et al. 2010)



The WMAP Haze

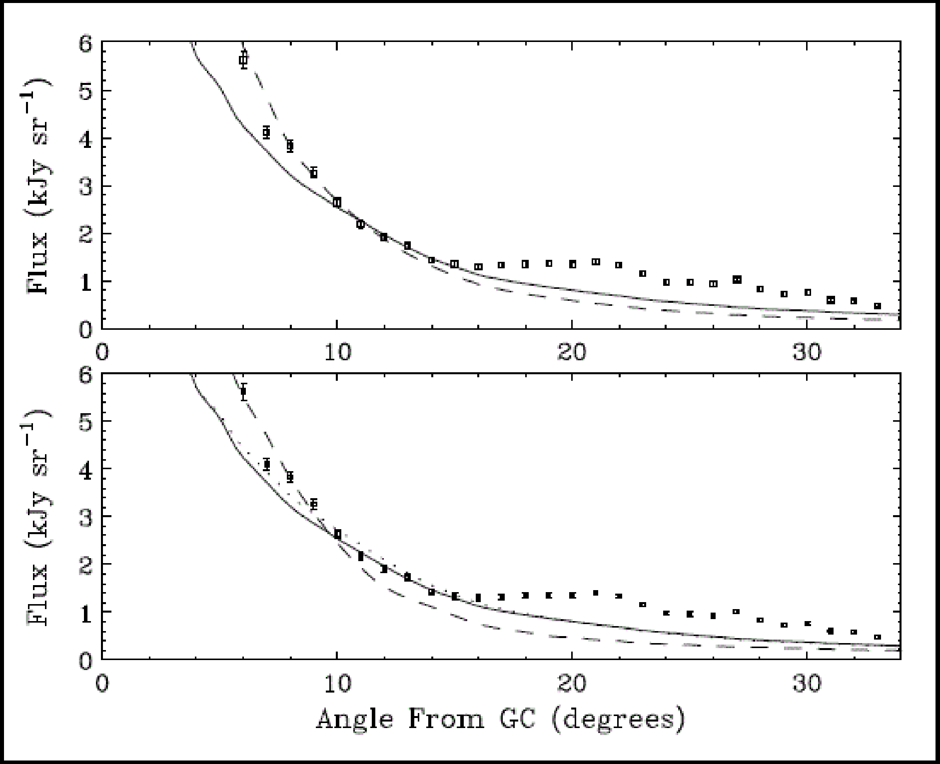
- Finkbeiner (2004) pointed out an unexplained residual in the WMAP dataset
- The existence of this residual is controversial, and is not detected by the WMAP team (Gold et al. 2010)



Hooper et al. (2007) (0705.3655)

The WMAP Haze

- Hooper et al. (2007) explained the WMAP haze as the result of dark matter annihilation
- Also explained by pulsars (Kaplinghat et al. 2009)

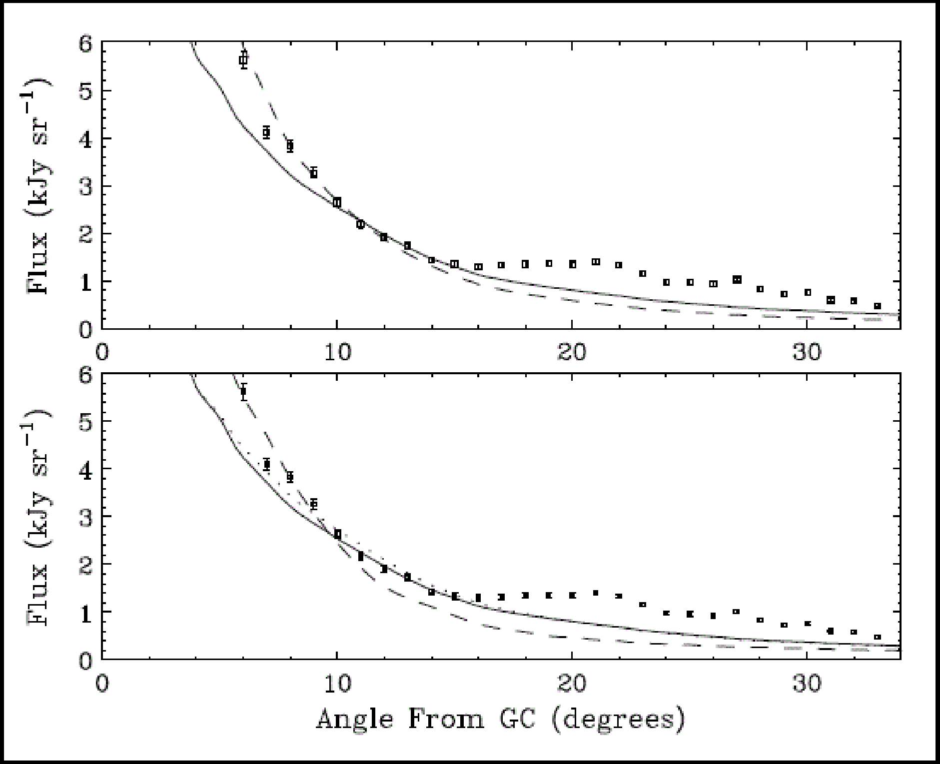


Hooper et al. (2007) (0705.3655)

The WMAP Haze

- The dark matter matches to this haze depended on non-standard diffusion parameters

- M_X = 100 GeV
- B = 10 μG
- XX → e^+e^-
- NFW Profile
- $D_0 = 1.58 \times 10^{28}$ $cm^2 s^{-1}$ (4 GeV)



Hooper et al. (2007) (0705.3655)

Research Goals

- Evaluate a select range of well motivated WIMP theories

- Test the DM interpretation of the WMAP haze using cosmic ray propagation models that are consistent with all current observations and data

Our modeling code

- Use DarkSUSY to calculate the primary e^+e^- spectrum for a range of well motivated DM models
- Use Galprop to determine the synchrotron emission and nuclear abundances in each propagation model
- Isolate the simulated DM haze by subtracting the synchrotron component from the corresponding simulation with DM disabled.

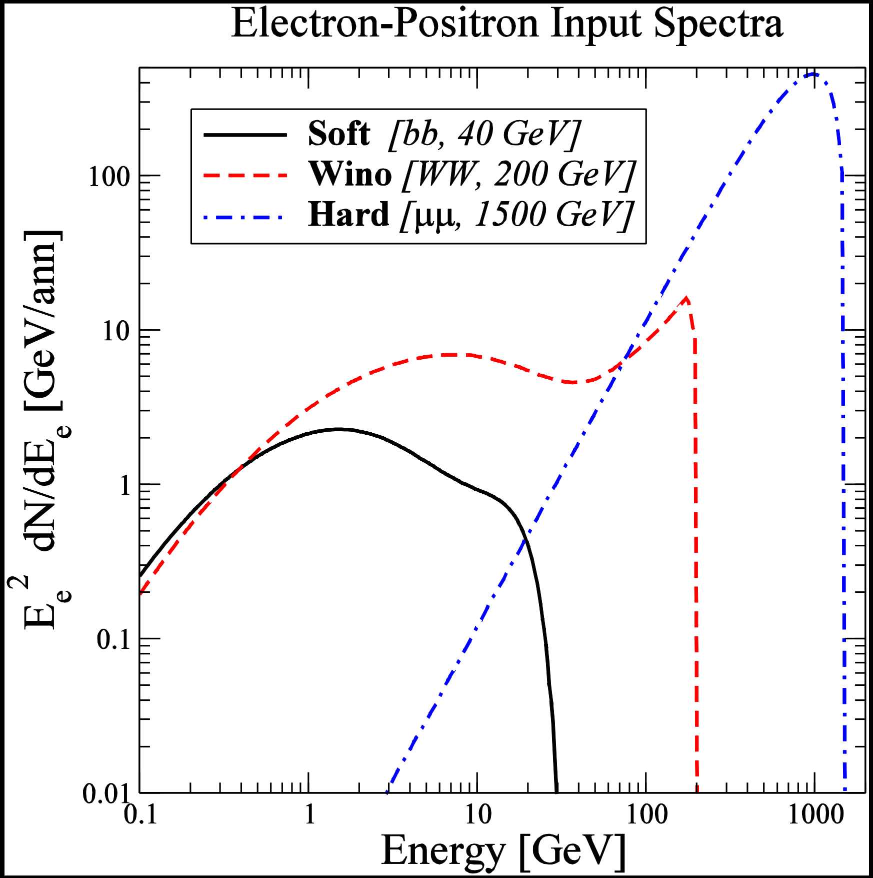
Dark Matter Models

- We test three DM annihilation channels which span a range of motivated WIMP decay models

Soft (40 GeV XX → b b-bar)

Wino (200 GeV XX → W^+W^-)

Hard (1500 GeV XX → $\mu^+\mu^-$)



Galprop Models

- Employing the public version of Galprop, we use the following parameters in our default setup:
  - $D_0 = 5.0 \times 10^{28}$ cm^2 s^{-1}
  - Simulation Height = 4 kpc
  - $V_{alfven} = 25$ km s^{-1}
  - Convection = Disabled
  - $B = 11.6 \exp(-r / 10\text{kpc} - z / 2\text{kpc})$ μG

Boost Factors

- We multiply the simulated haze by a universal constant to match the observed WMAP haze at 10 degrees latitude and 23 Ghz.

$\Phi = \rho^2(x)/M_{DM}^2 \ \langle\sigma v\rangle$ $\langle\sigma v\rangle \sim 3 \times 10^{-26}\ cm^2 s^{-1}$

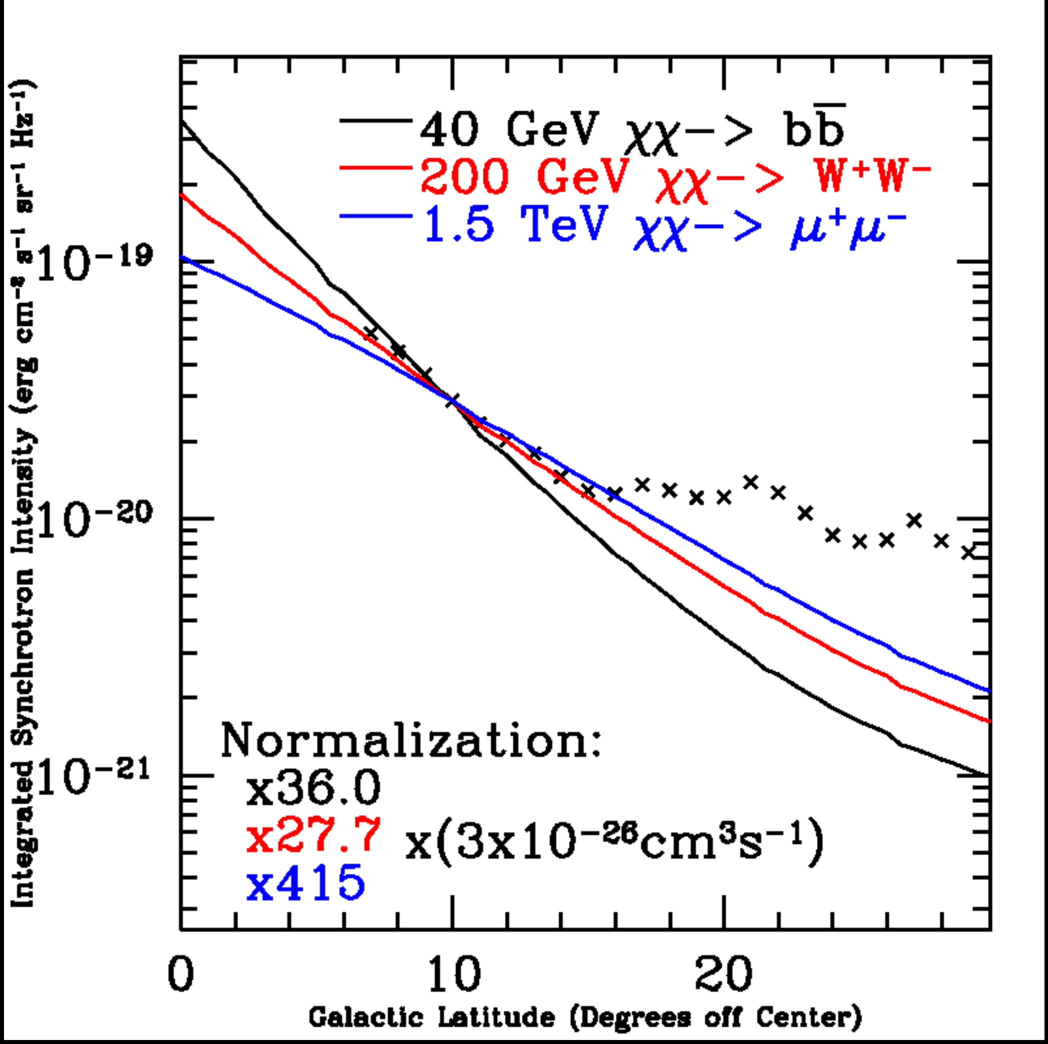
- Changes in $\langle\sigma v\rangle$

- Density fluctuations in DM substructure

- Sommerfield enhancements

Default Model Predictions

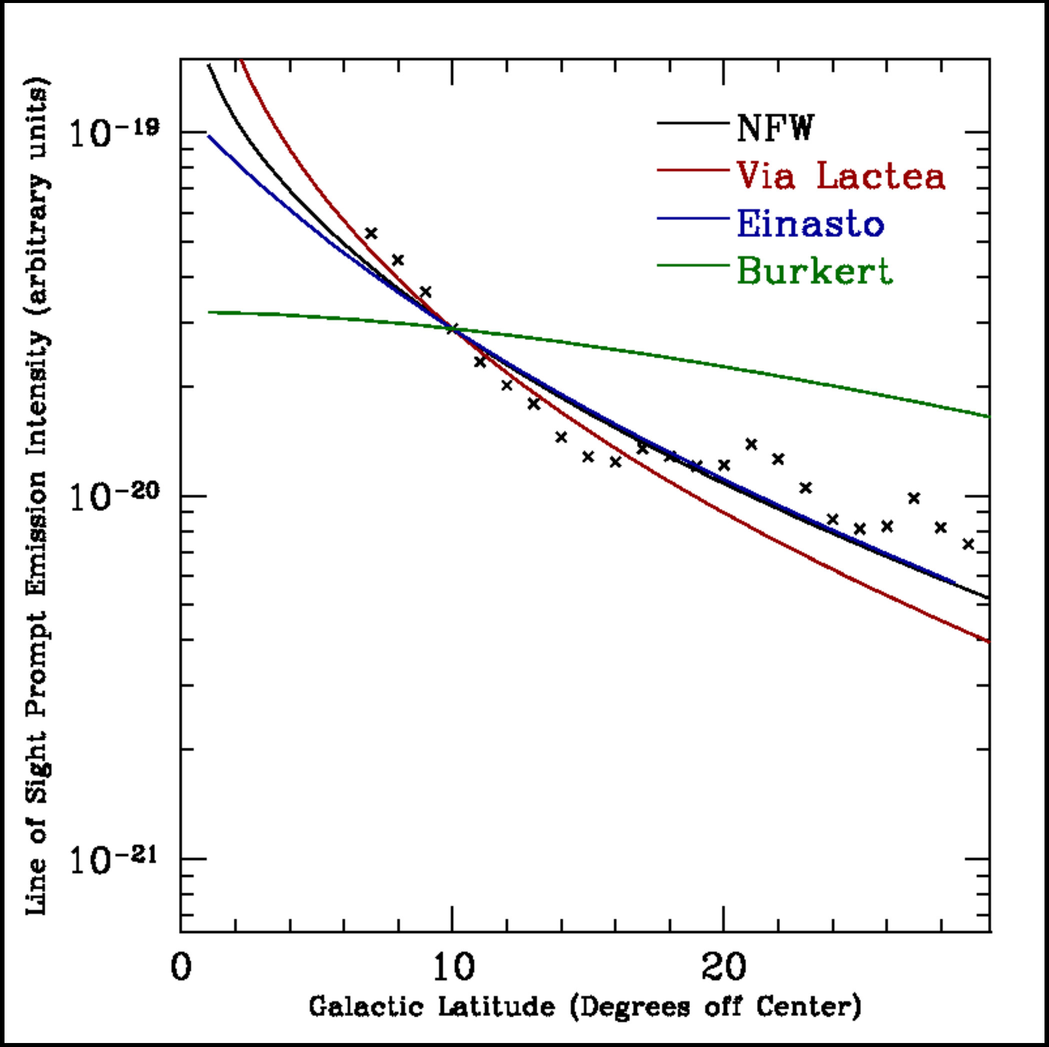
- Our default model shows a morphology which falls off much faster as a function of latitude than the observed haze
- Similar underestimates as in Fermi Haze?
- The WMAP haze requires large boost factors



Role of Diffusion

- Without diffusion, the DM profiles actually suggest a much flatter distribution

- Diffusion plays <u>counterintuitive</u> role of increasing the falloff in emission at high latitudes



Role of Diffusion

- We test four diffusion parameters:
  - Diffusion Constant (D_0)
    - Ability of charged particles to move through galaxy
    - Can be thought of as the "thickness" of the soup the particles move through
  - Simulation height (z)
  - Alfvén Velocity (v_α)
  - Convection Velocity

Role of Diffusion

- We test four diffusion parameters:
  - Diffusion Constant (D_0)
  - Simulation height (z)
    - Height of zone which particles move through before they exit the "soup" of the galaxy
  - Alfvén Velocity (v_α)
  - Convection Velocity

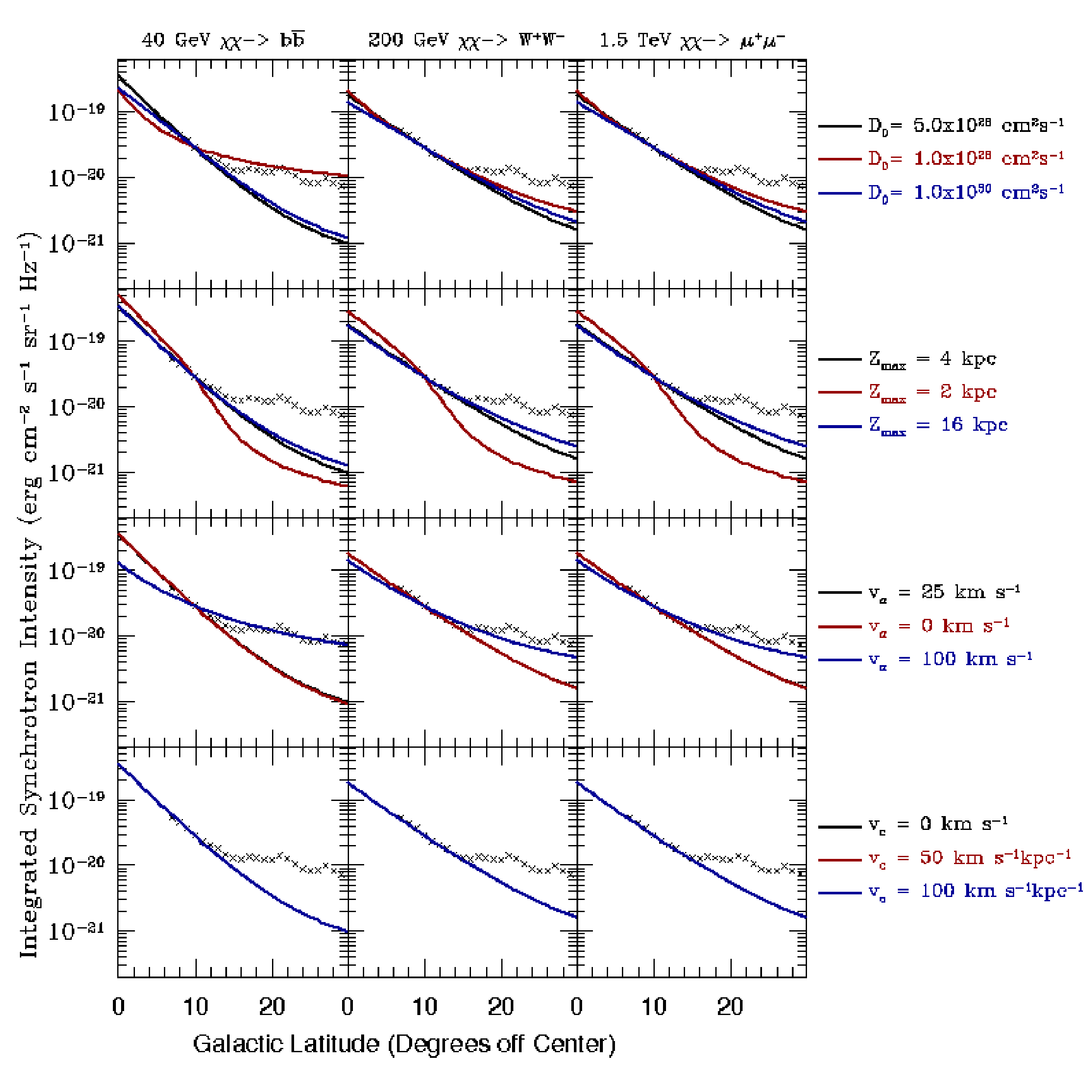
Role of Diffusion

- We test four diffusion parameters:
  - Diffusion Constant (D_0)
  - Simulation height (z)
  - Alfvén Velocity (v_α)
    - Diffusion of particles through momentum space
    - Reacceleration of particles
  - Convection Velocity

Role of Diffusion

- We test four diffusion parameters:
  - Diffusion Constant (D_0)
  - Simulation height (z)
  - Alfvén Velocity (v_α)
  - Convection Velocity
    - Cosmic “wind” pushing particles out of the galaxy

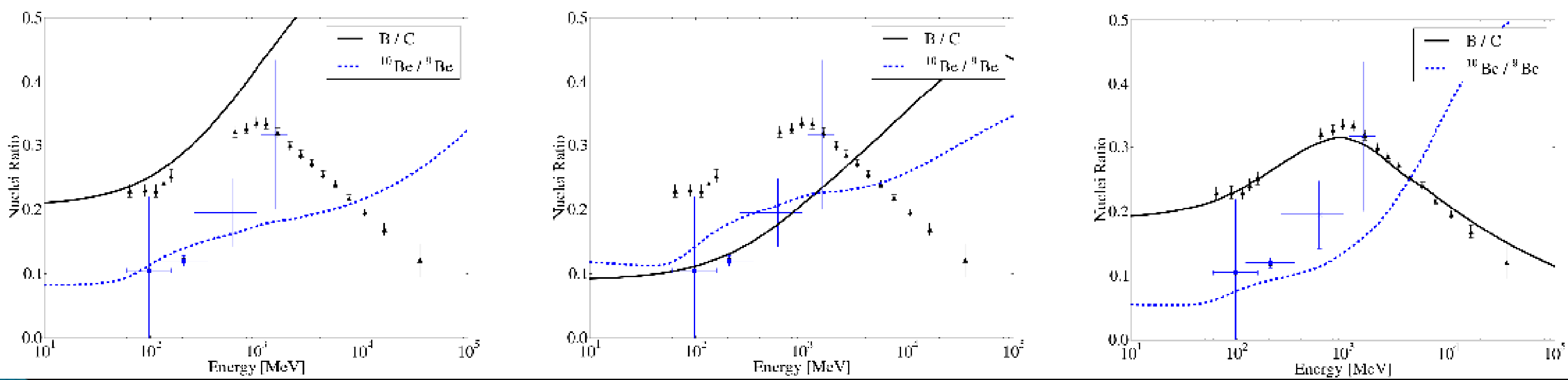
Role of Diffusion



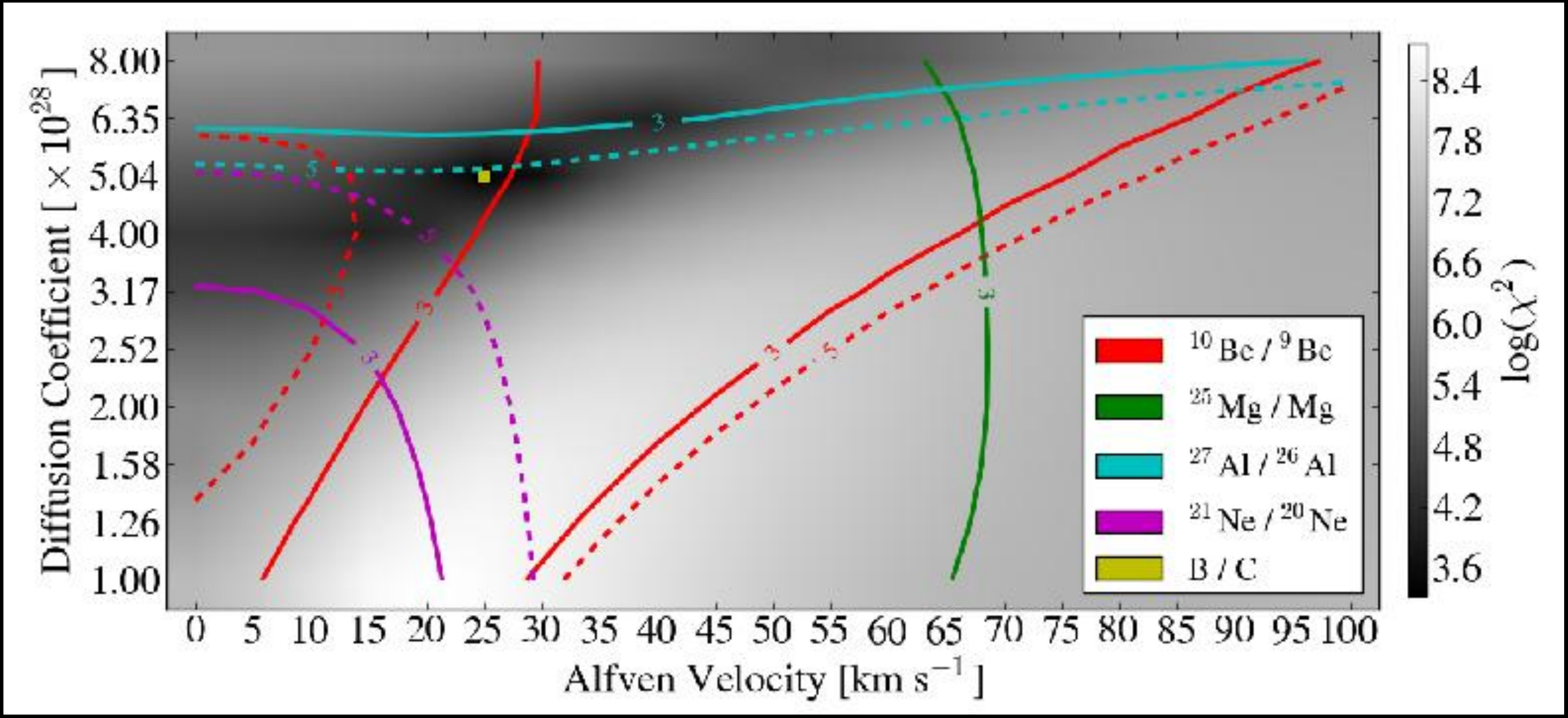
- The diffusion constant and Alfvén velocity greatly affect the Haze morphology

Constraints on Diffusion

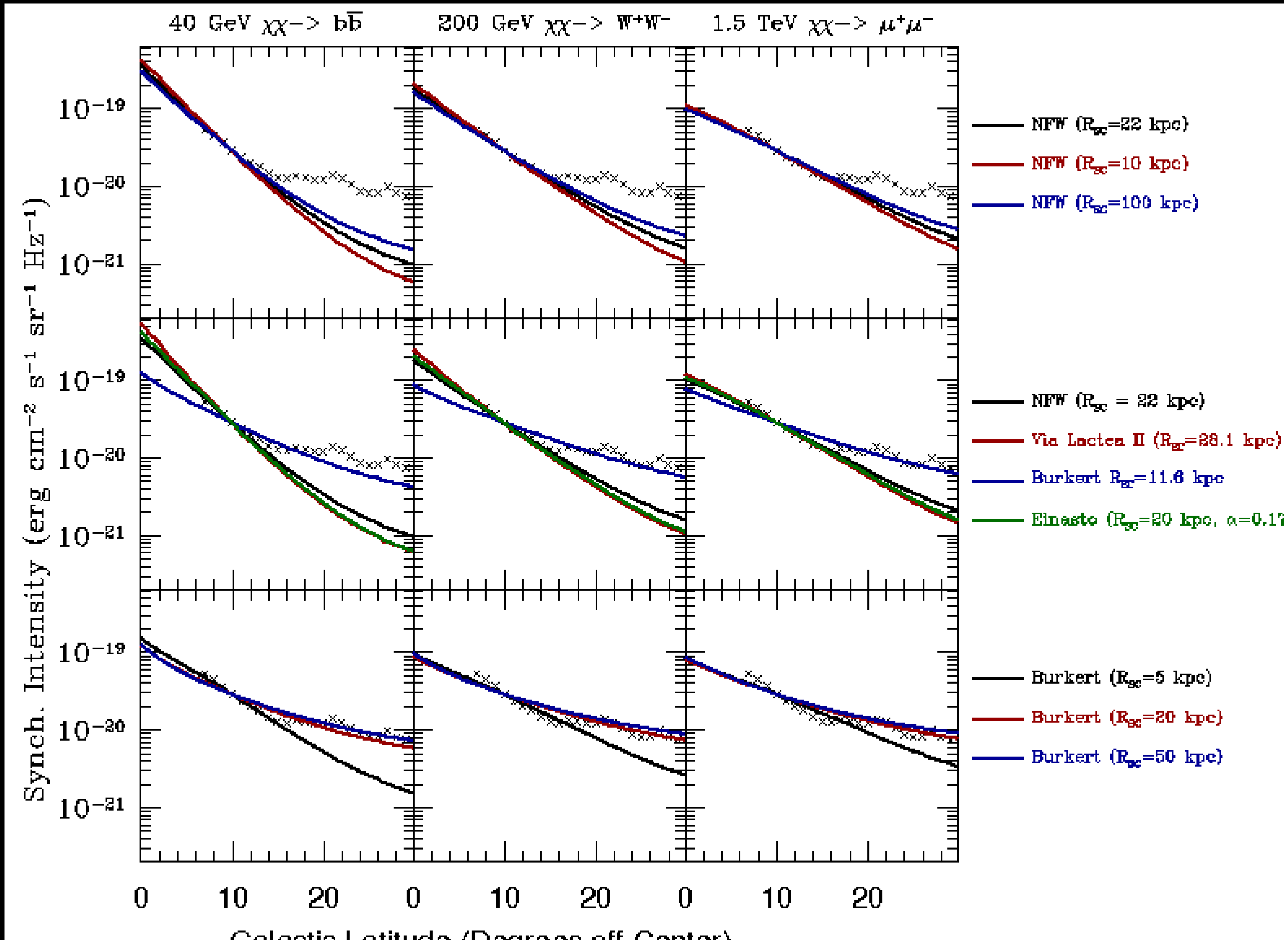
- Changes in the diffusion setup will affect the ratio of cosmic ray primary to secondary species
- This allows changes in the diffusion setup to be constrained by local cosmic ray observations



Constraints on Diffusion



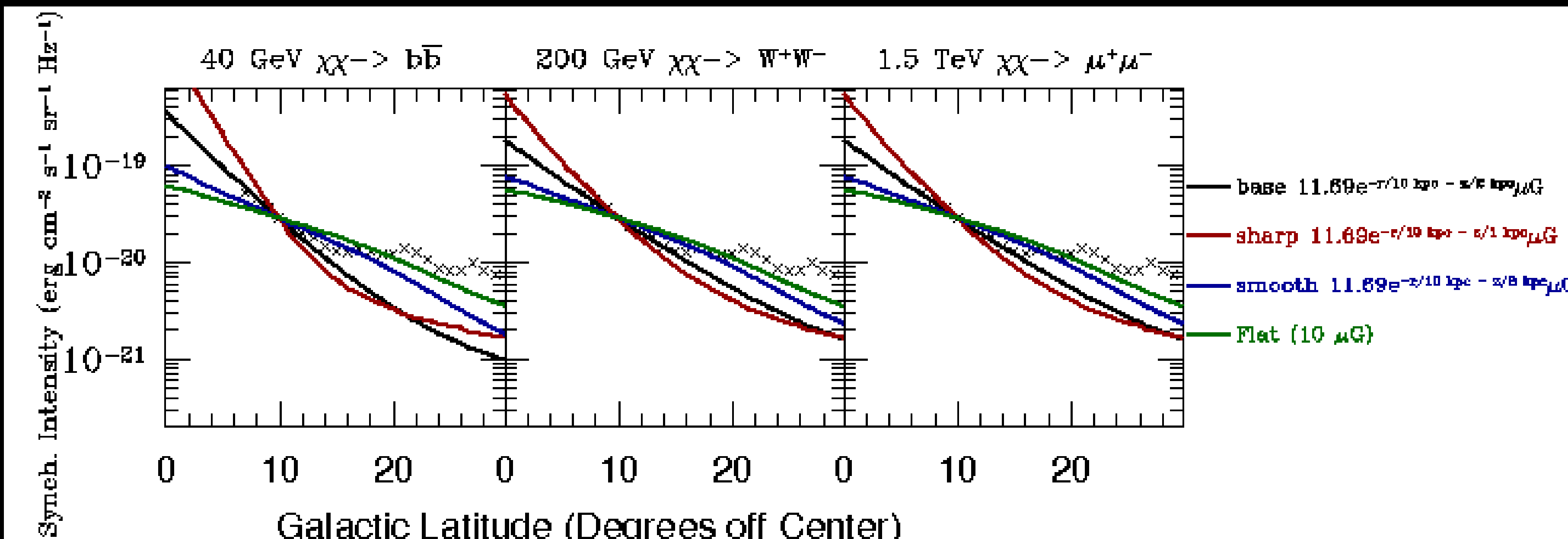
DM Profiles



Only profile which brings a reasonable match to the WMAP haze is a Burkert profile

- Profiles with dense galactic centers are unable to recreate the haze

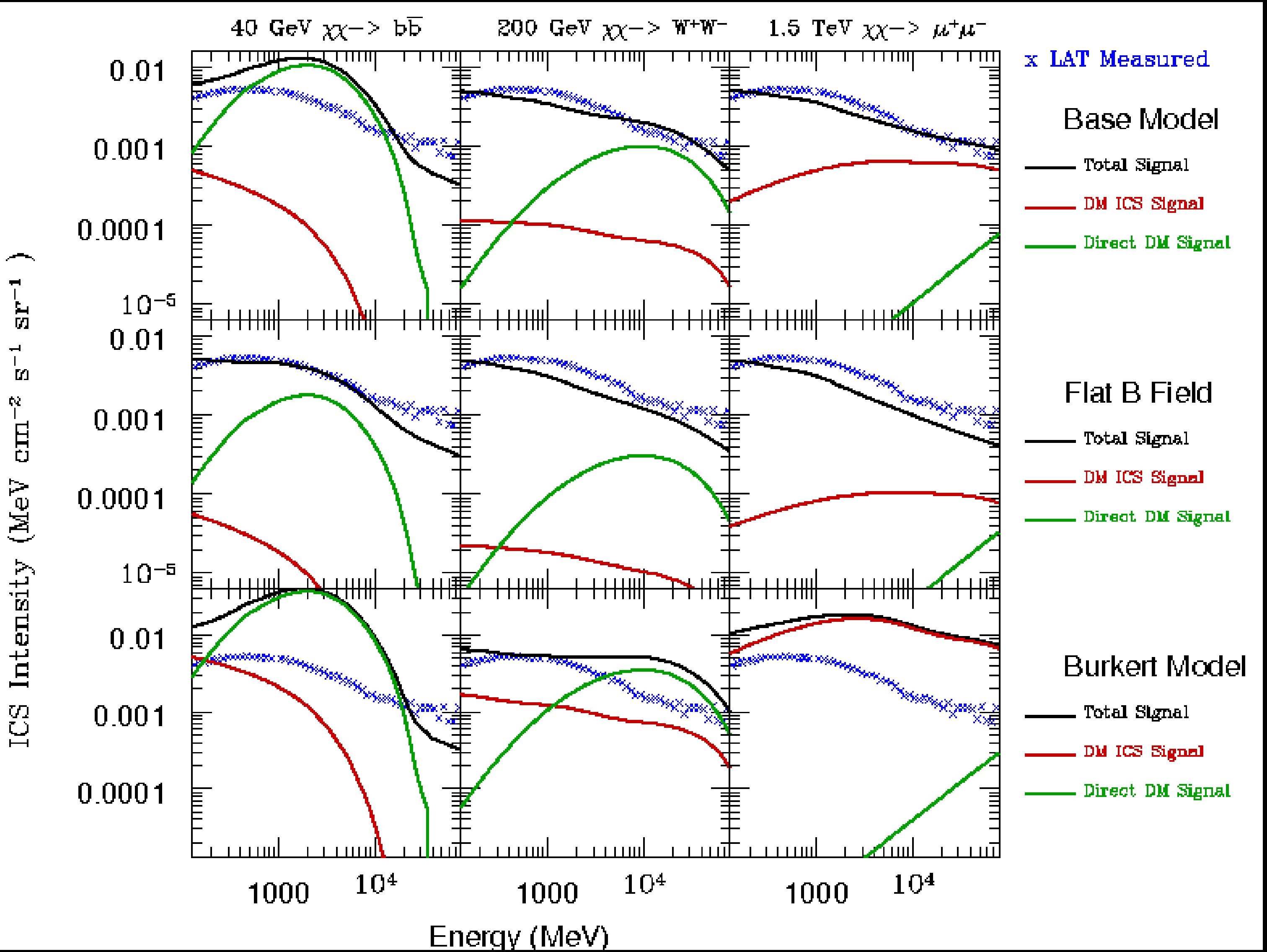
Magnetic Field Models



- Magnetic fields are an important uncertainty in our models

WMAP Matches?

- We have two possible matches to the morphology of the WMAP haze:
  - Changes in the magnetic field distribution (Flat magnetic field)
  - Changes in the DM density distribution (Burkert profile)
- Changes in the diffusion parameters have been ruled out by cosmic ray constraints



ICS

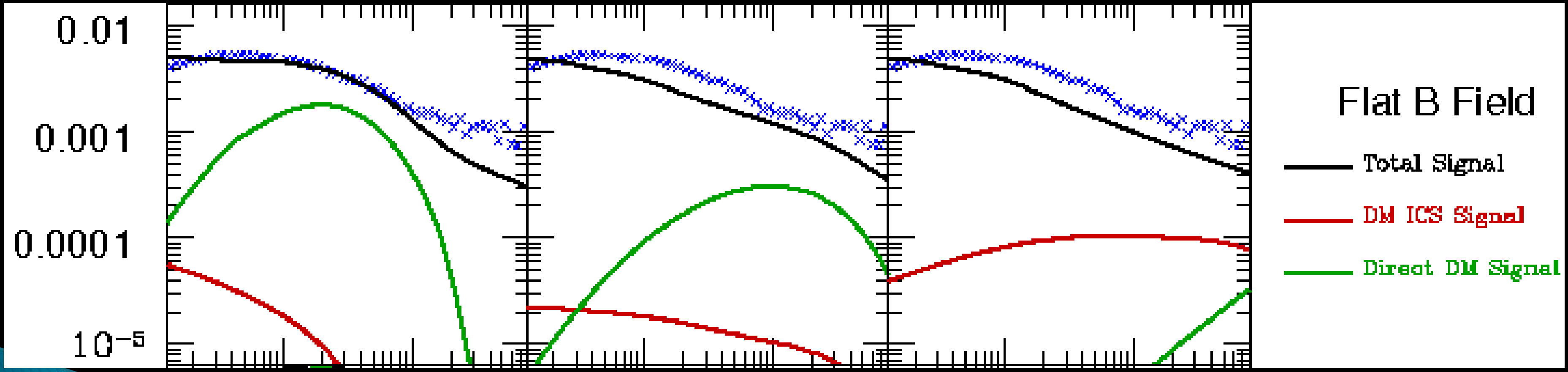
- Expected Fermi signals from our "matching" profiles

Conclusions

- Standard Dark Matter/Diffusion setups do not provide a reasonable match to the WMAP haze
- Diffusion setups that would match the WMAP haze are well constrained by cosmic ray observations
- DM profiles which would move annihilations to higher latitudes are well constrained by Fermi observations
- Magnetic field models are a major uncertainty

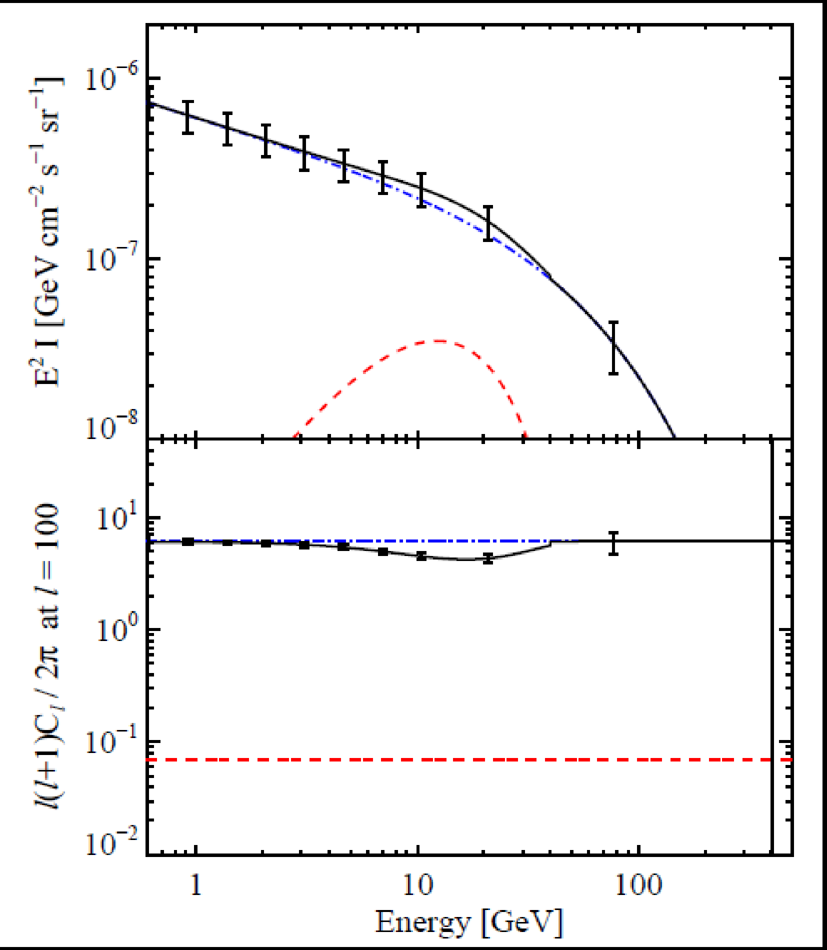
Anisotropy Studies

- Very large backgrounds can be hiding inside the Fermi signal
- Uncertainties in astrophysical backgrounds make difficult to distinguish between signals



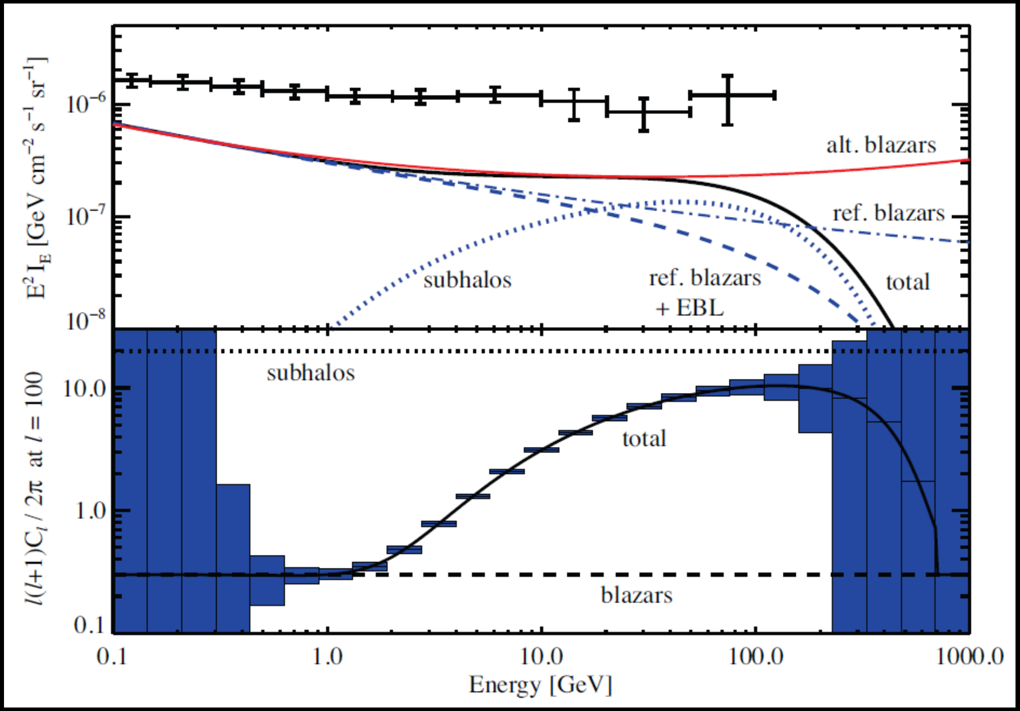
Anisotropy Studies

- Studying anisotropies allow us to pick small backgrounds out of large foreground signals



Hensley, Siegal-Gaskins, Pavlidou (2009)

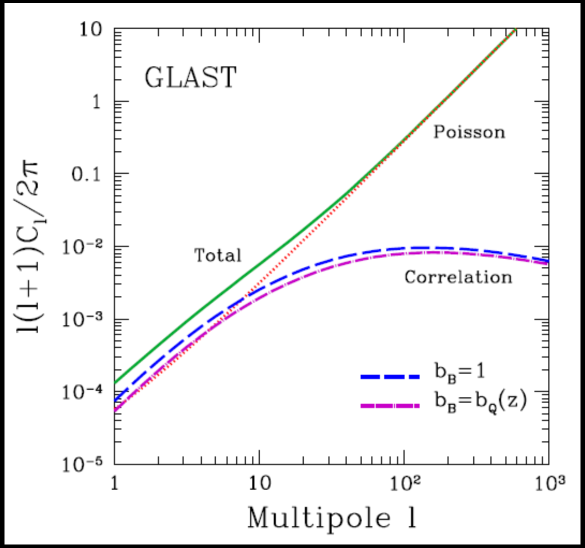
Anisotropy Studies



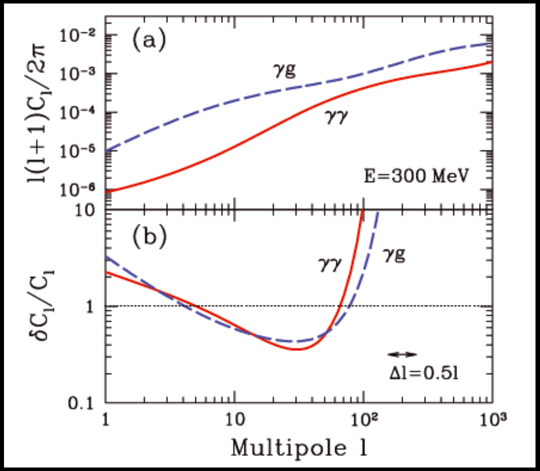
Siegal-Gaskins and Pavlidou (2009)

- Measurements of the anisotropy as a function of energy allow us to differentiate between different source classes

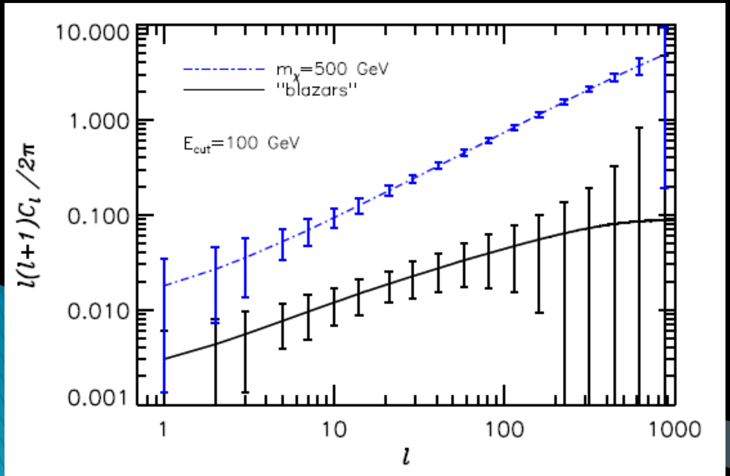
Anisotropy Studies



Blazars – Ando et al (2007)



Starforming Galaxies – Ando and Pavlidou (2009)



Extragalactic DM
Cuoco et al. (2008)

Conclusions

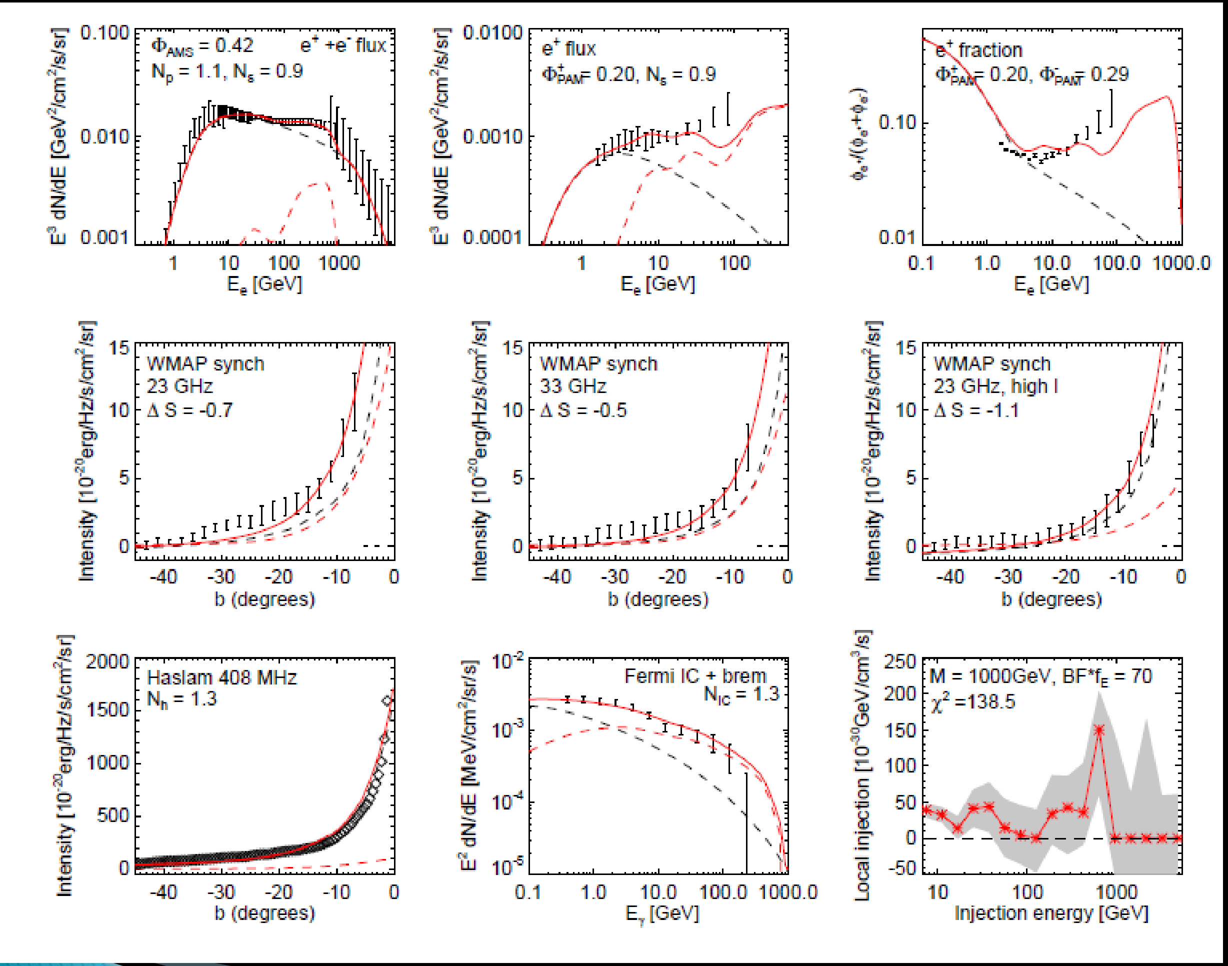
- 1.) While the spectrum and intensity of the Fermi haze is systematically uncertain due to foreground templates, there does appear to be an anomolous foreground

- 2.) Uncertainties in astrophysical templates make it very difficult to understand these anomolies

- 3.) New methods will be necessary to understand these observations

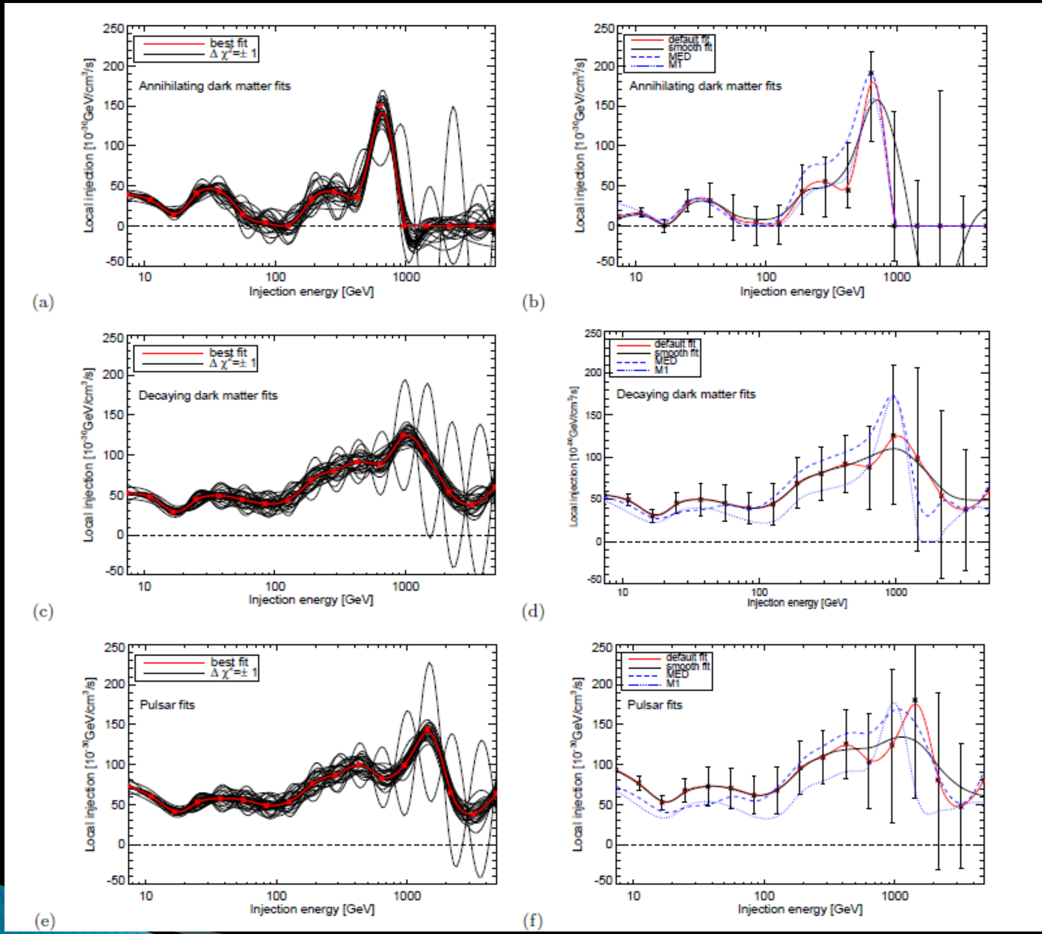
EXTRA SLIDES

Matching the Fermi Haze



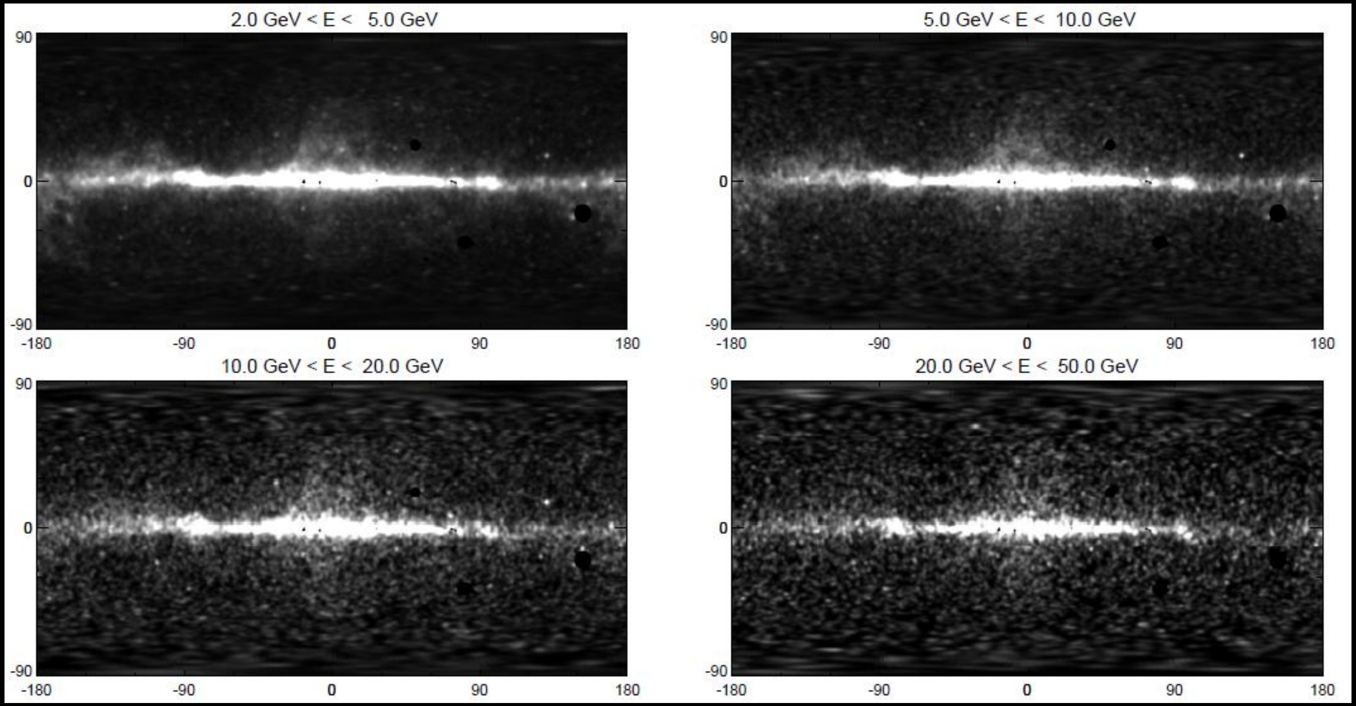
Liu et al. (2010)

Matching the Fermi Haze



Liu et al. (2010)

The Fermi Sky

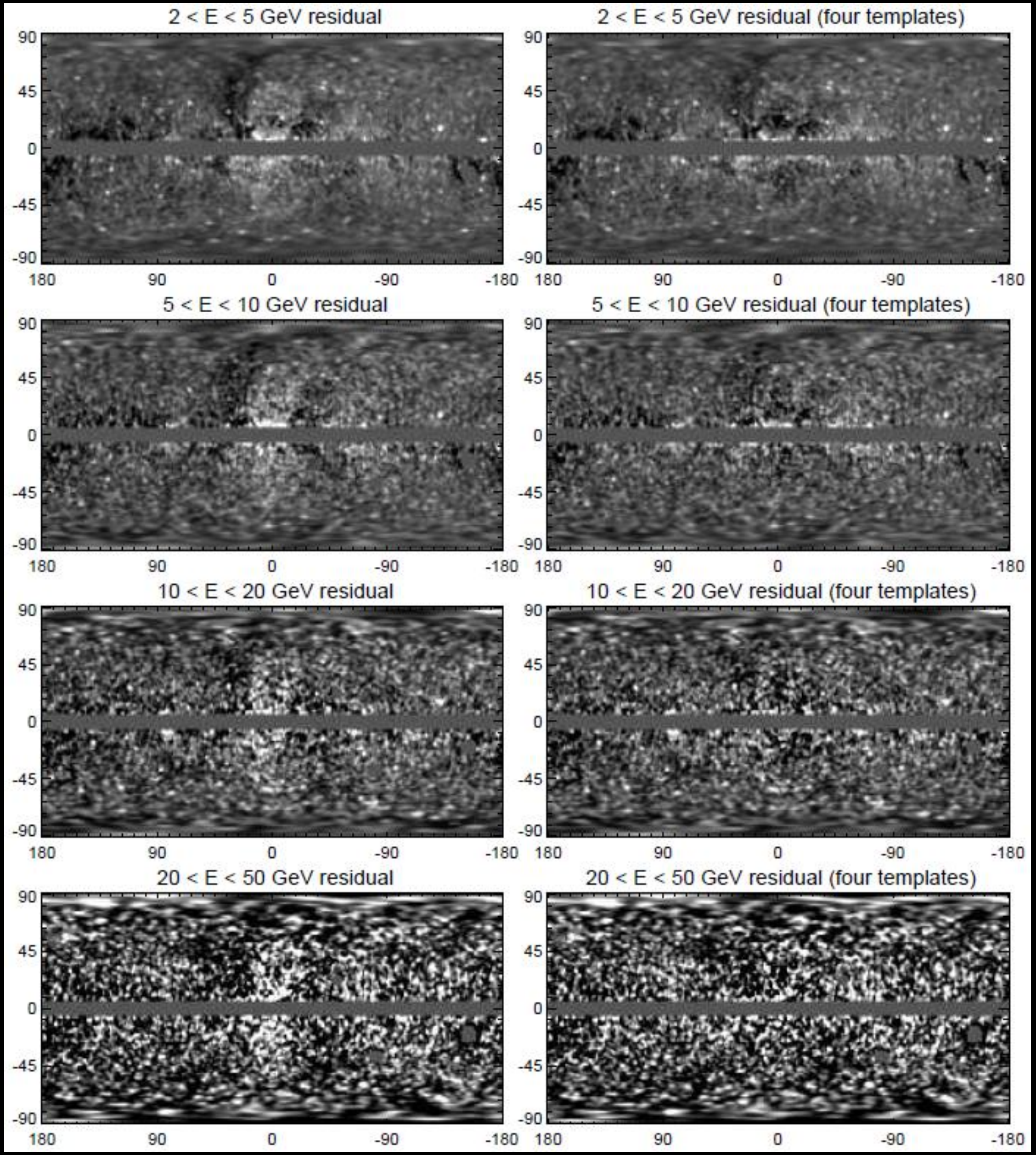


Dobler et al (2009) produced skymaps of the Fermi diffuse emission

- This skymaps come from many classes –

Results

- Haze template effectively eliminates the morphology of the residual
- Residual now appears to be random noise



Matching this spectrum

- Nevertheless, it is worth testing whether we can match this spectrum with new electron inputs (specifically dark matter)

- We employ Galprop models including a dark matter contribution, and determine the morphology of the output WMAP spectrum

Matching this spectrum

- Nevertheless, it is worth testing whether we can match this spectrum with new electron inputs (specifically dark matter)
- We employ Galprop models including a dark matter contribution, and determine the morphology of the output WMAP spectrum